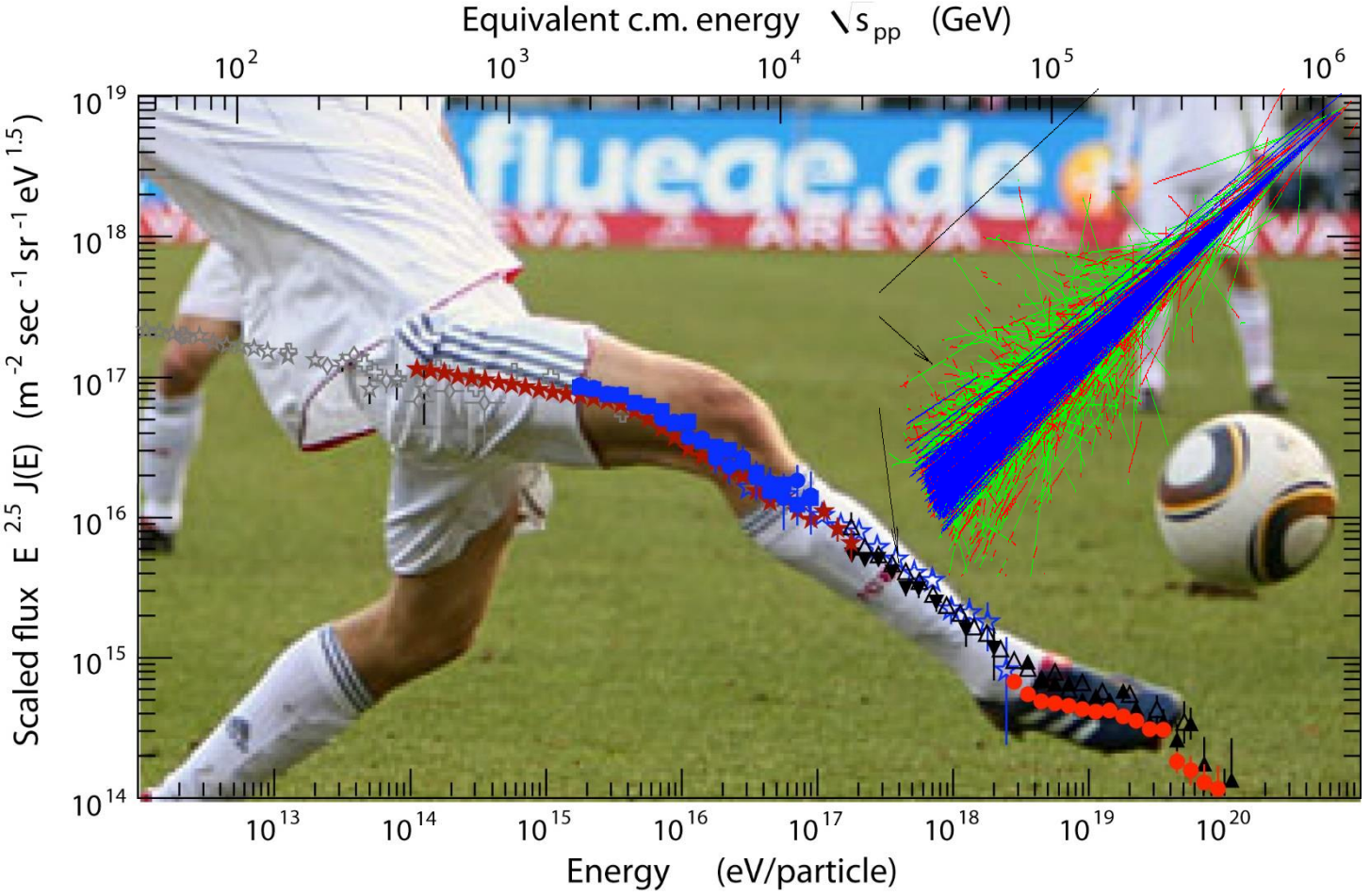
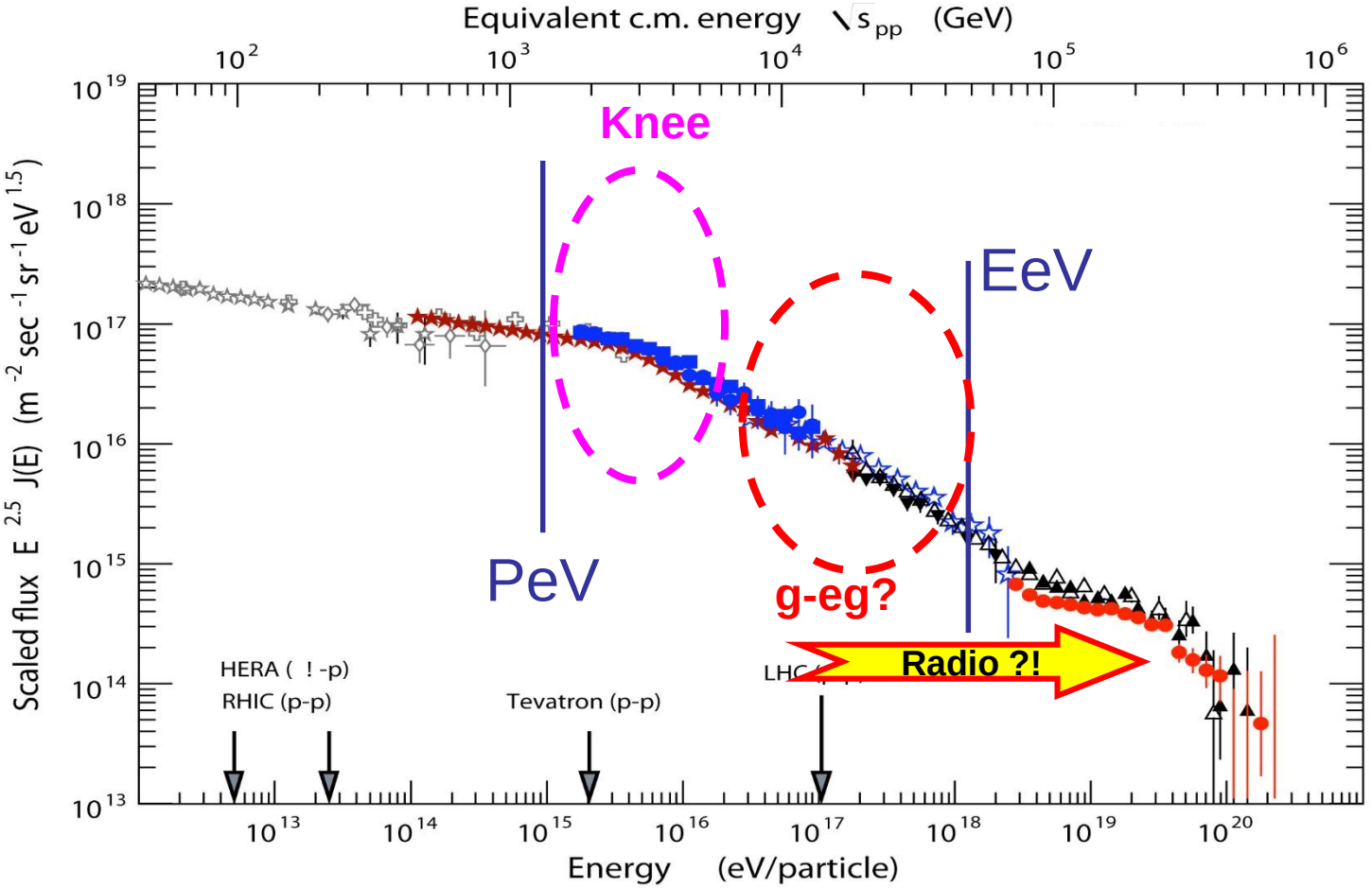
KASCADE-Grande

in the view of the post-LHC hadronic interaction models

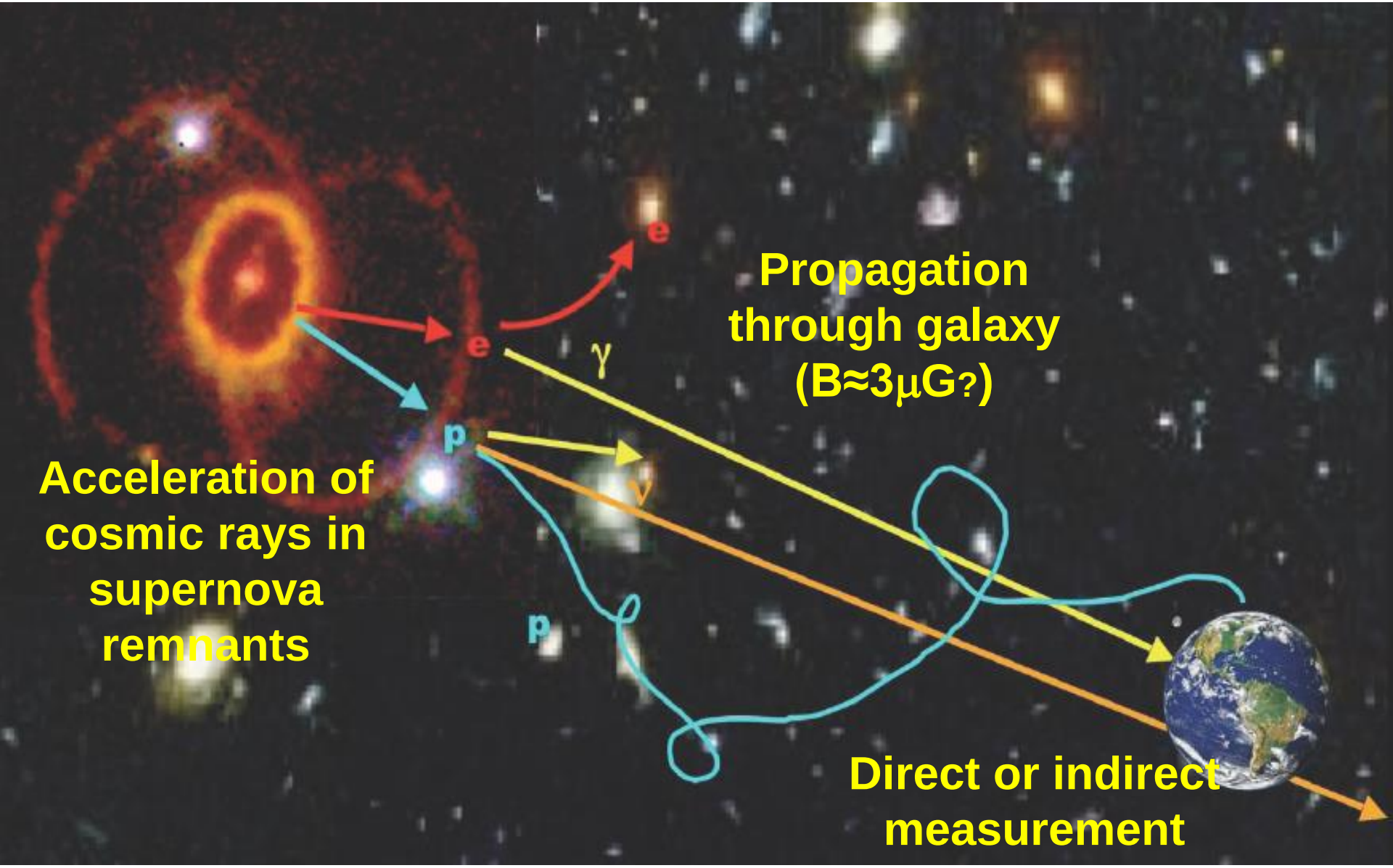


The physics of the KASCADE facility



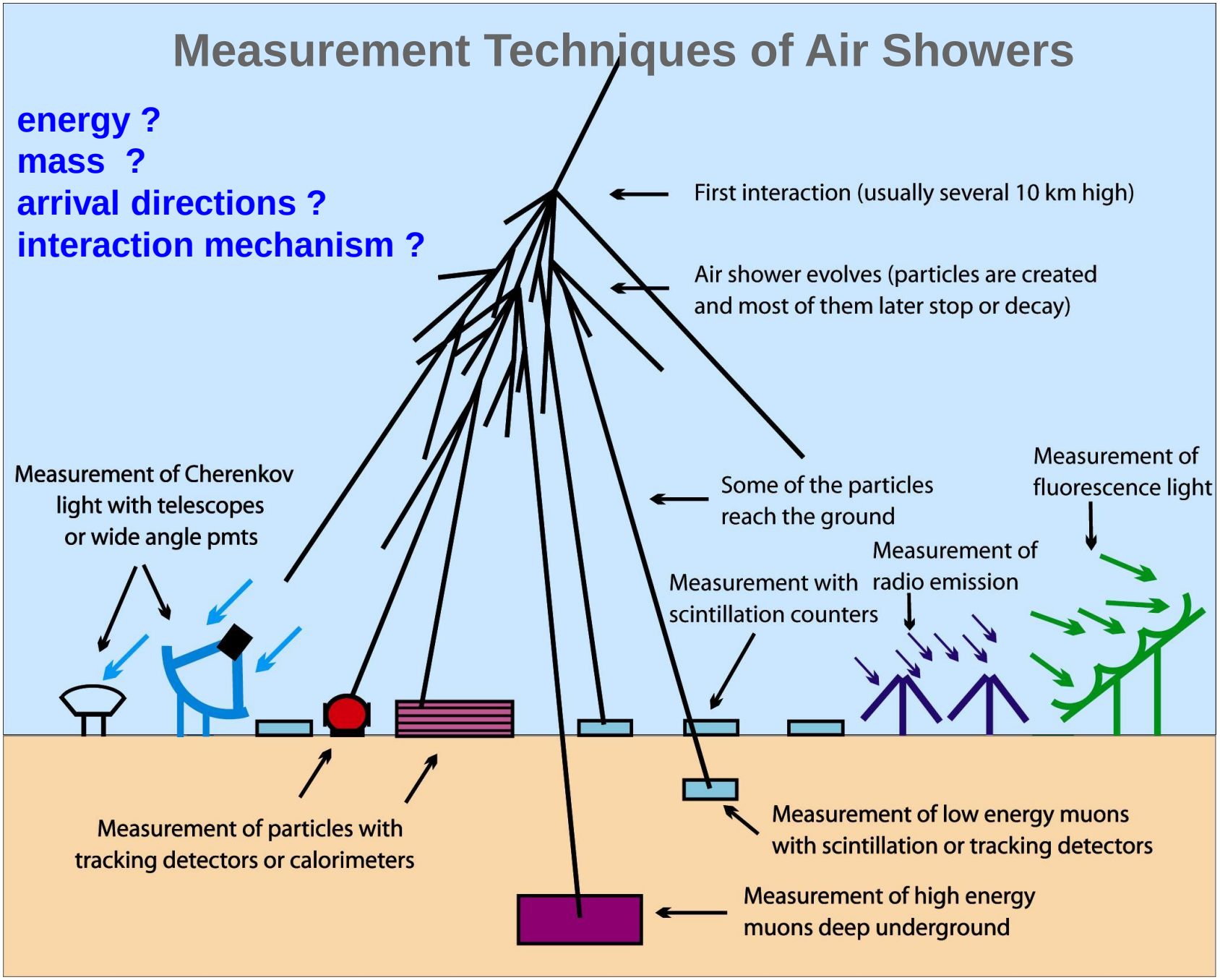
KASCADE	***10^{15}-10^{17}eV:***	**Origin of the knee?**
KASCADE-Grande	***10^{16}-10^{18}eV:***	**Iron knee (rigidity)?**
		Transition galactic-eg CR?
LOPES	***$10^{16.7}$-10^{18}eV:***	**New detection technique!**

Galactic cosmic rays



Measurement Techniques of Air Showers

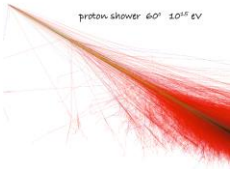
energy ?
mass ?
arrival directions ?
interaction mechanism ?



KASCADE timeline

- 55 collaborative papers in reviewed journals (4-5 still in queue, short author list papers not included)
- 56 PhD thesis
- 86 diploma/master thesis

1989 Proposal
CORSIKA
1996
KASCADE
2000
model tests
light knee
2004
γ limit
(an)isotropy
proof-of-principle radio detection
muon production height
2008
heavy knee
light ankle
EAS GHz emission
Xmax by radio
2012
KCDC
CROME
Karlsruhe Air Shower Test Facility
Cosmic Revelation
LOPES
KASCADE-Grande
2016



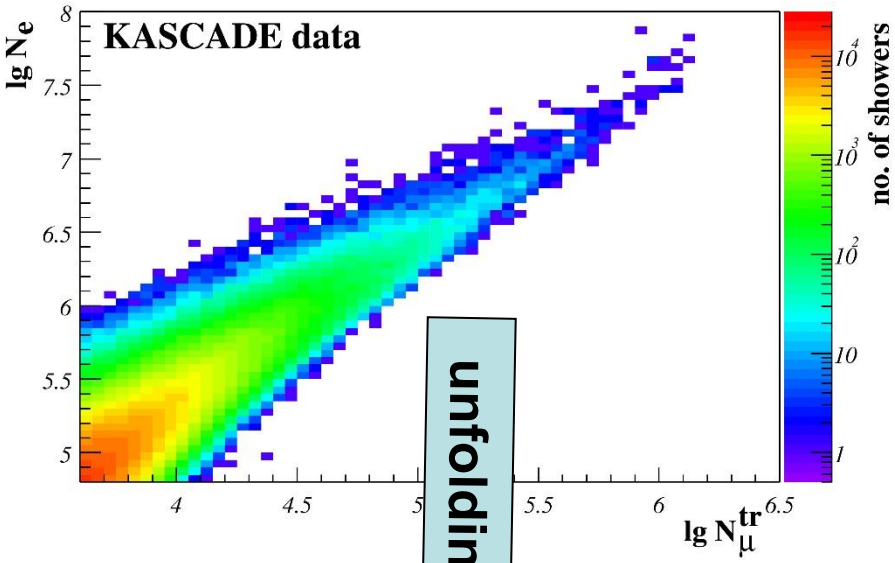
KASCADE

KArlsruhe Shower Core and Array DEtector

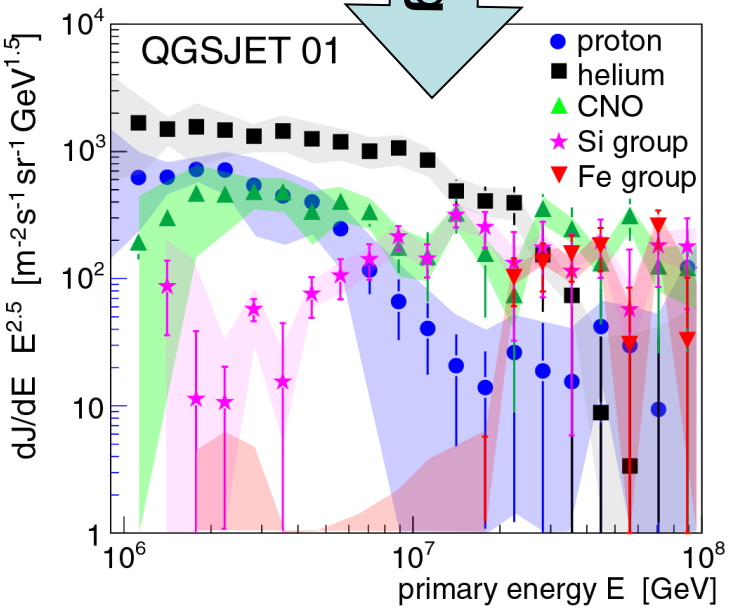
- Energy range 100TeV – 80PeV
- Since 1995
- Large number of observables: electrons, muons@4 thresholds, hadrons

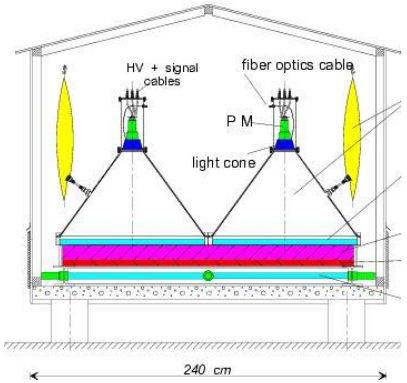
T.Antoni et al. NIM A513 (2003) 490

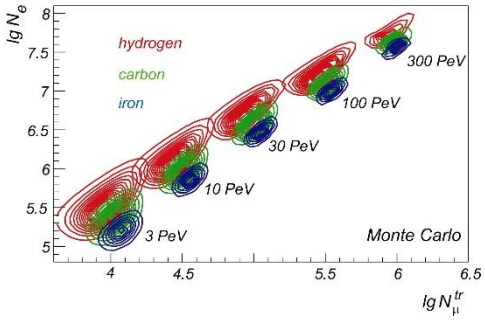
KASCADE : energy spectra of single mass groups



unfolding







Searched:

E and A of the Cosmic Ray Particles

Given:

N_e and N_μ for each single event

➔ solve the inverse problem

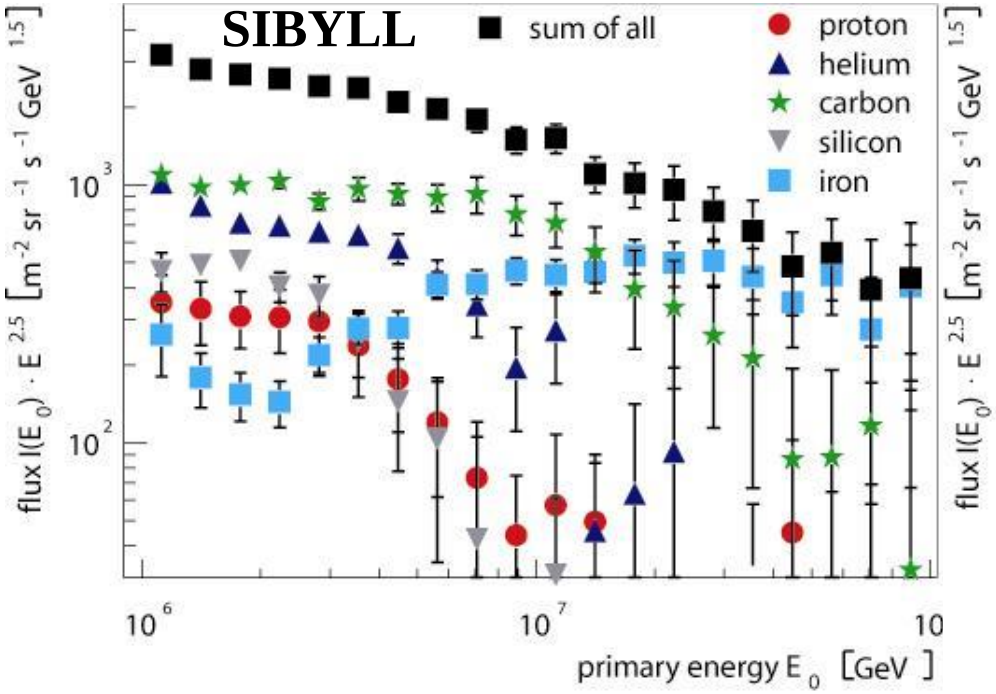
$$\frac{dJ}{d\lg N_e\, d\lg N_\mu^{tr}} = \sum_A \int_{-\infty}^{+\infty} \frac{dJ_A}{d\lg E}\, p_A(\lg N_e, \lg N_\mu^{tr} \mid \lg E)\, d\lg E$$

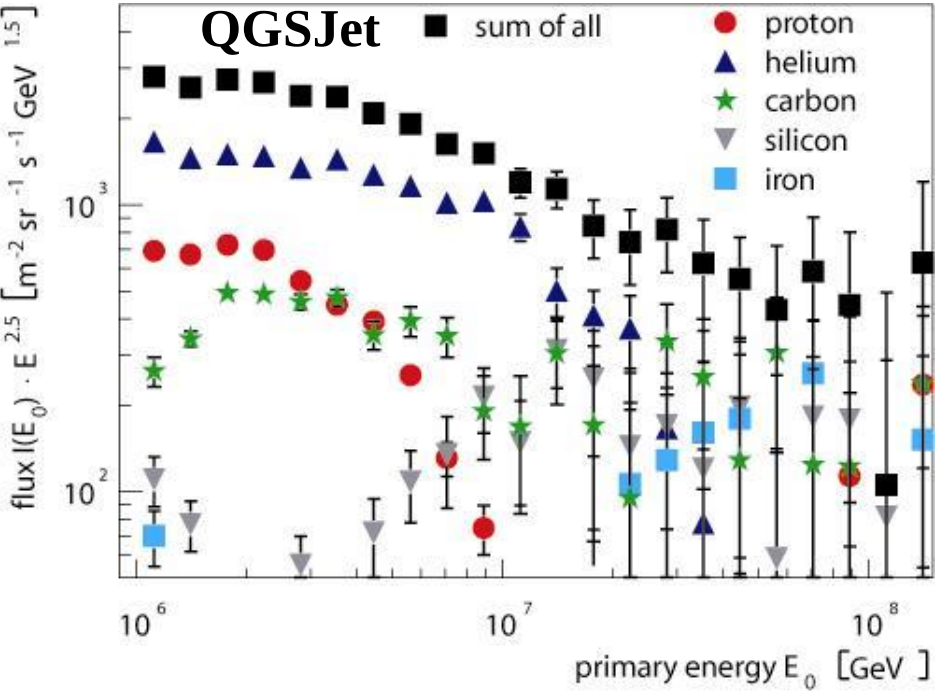
- kernel function obtained by Monte Carlo simulations (CORSIKA)
- contains: shower fluctuations, efficiencies, reconstruction resolution

KASCADE collaboration, Astroparticle Physics 24 (2005) 1-25

KASCADE: the rigidity knee

- same unfolding but based on different hadronic interaction models embedded in CORSIKA



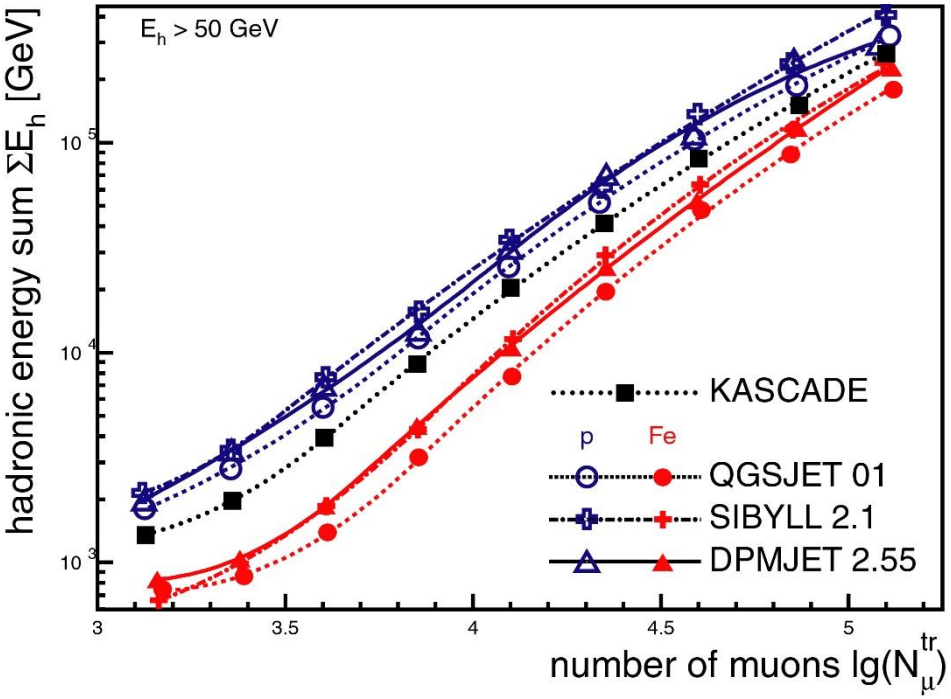


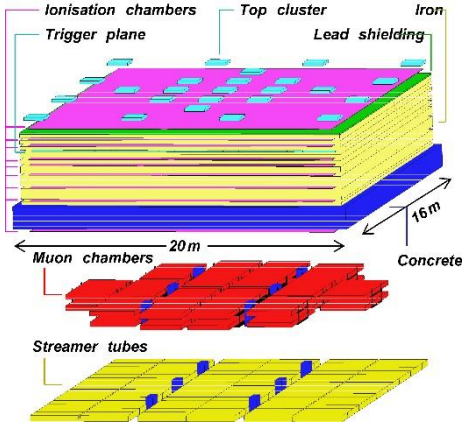
- all-particle spectrum similar
- general structure similar: knee by light component
- relative abundances very different for different high-energy hadronic interaction models

but for many models: proton not the most dominant component!

KASCADE collaboration, Astrop.Phys. 24 (2005) 1 , Astrop.Phys. 31 (2009) 86

KASCADE : sensitivity to hadronic interaction models





Example:
hadrons vs. muons

correlation of observables:
no hadronic interaction model describes data consistently !

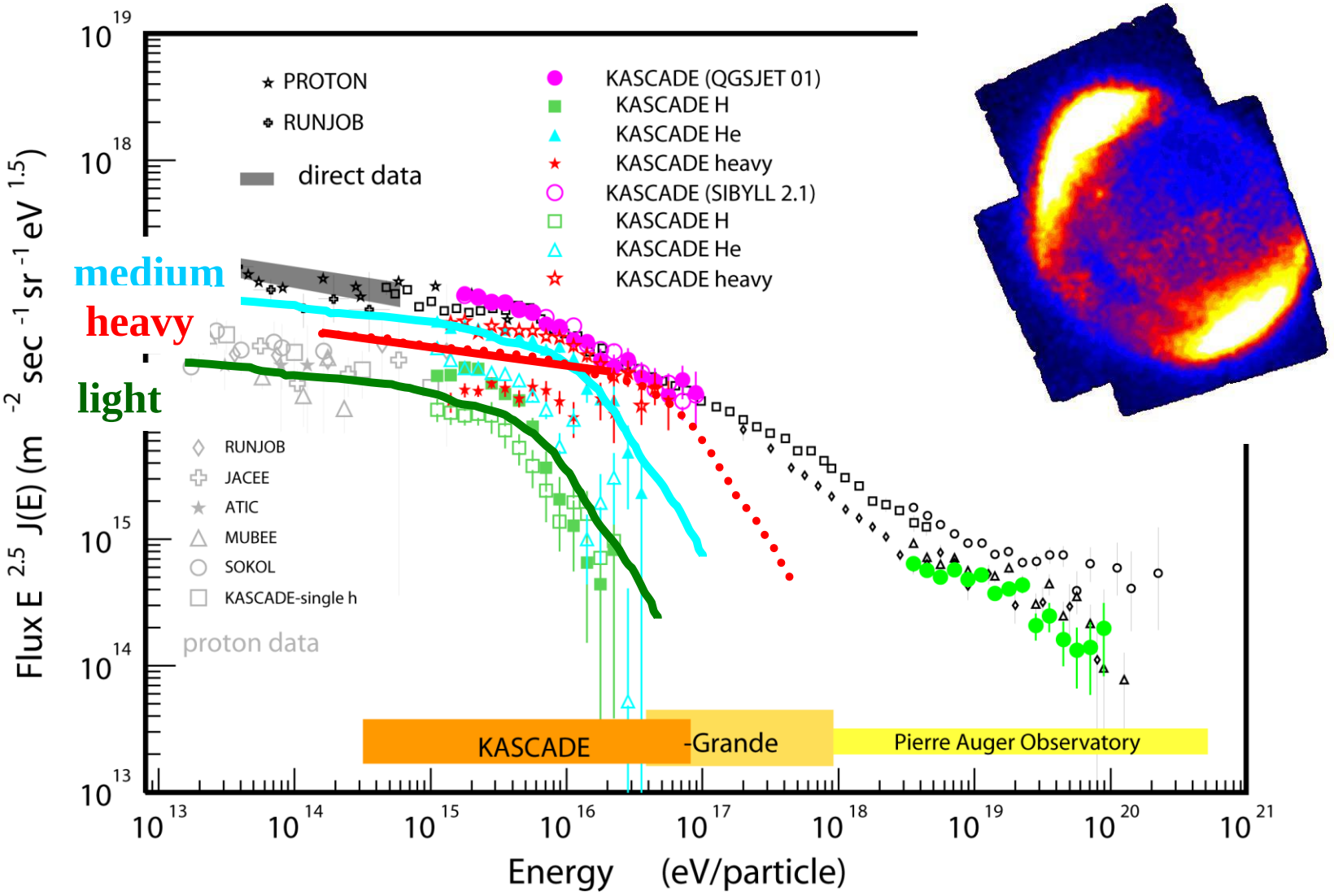
➔ **tests and tuning of hadronic interaction models !**

➔ **close co-operation with theoreticians (CORSIKA including interaction models)**

➔ **e.g.:**
- **EPOS 1.6 is not compatible with KASCADE measurements**
- **QGSJET 01and SIBYLL 2.1still most compatible models**

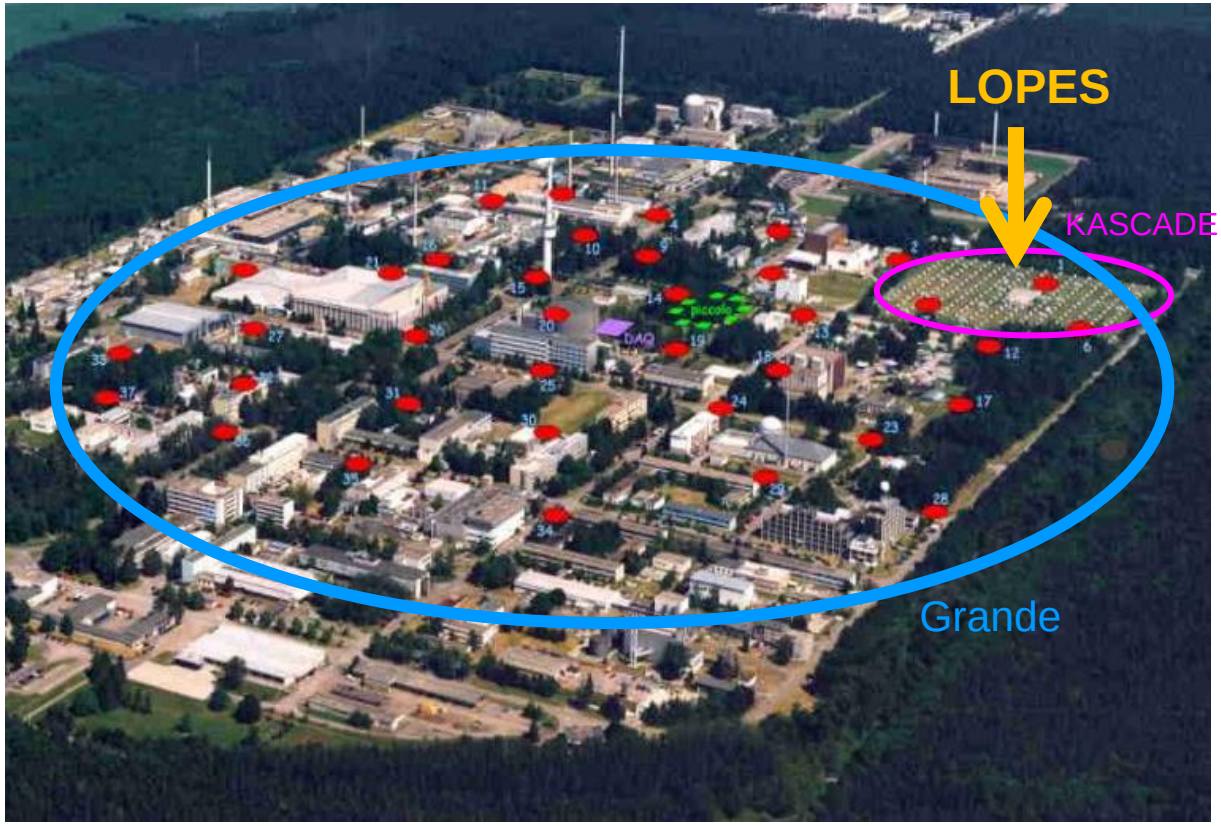
KASCADE collaboration, J Phys G (3 papers: 25(1999)2161; 34(2007)2581; (2009)035201)

Result KASCADE ➔ Motivation KASCADE-Grande



LOPES

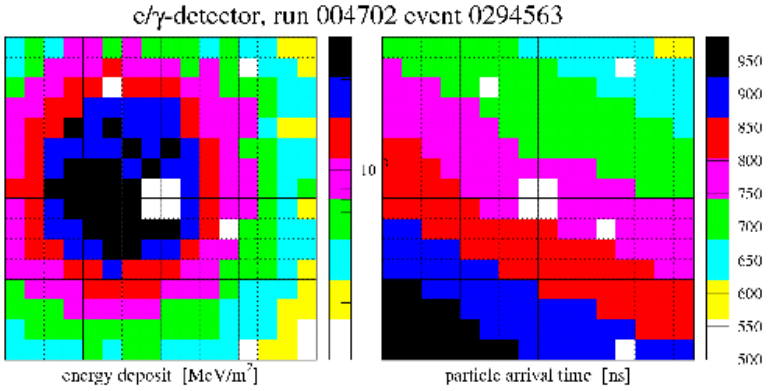
LOPES collaboration:
-) KASCADE-Grande
-) U Nijmegen, NL
-) MPIfR Bonn, D
-) Astron, NL
-) IPE, FZK, D



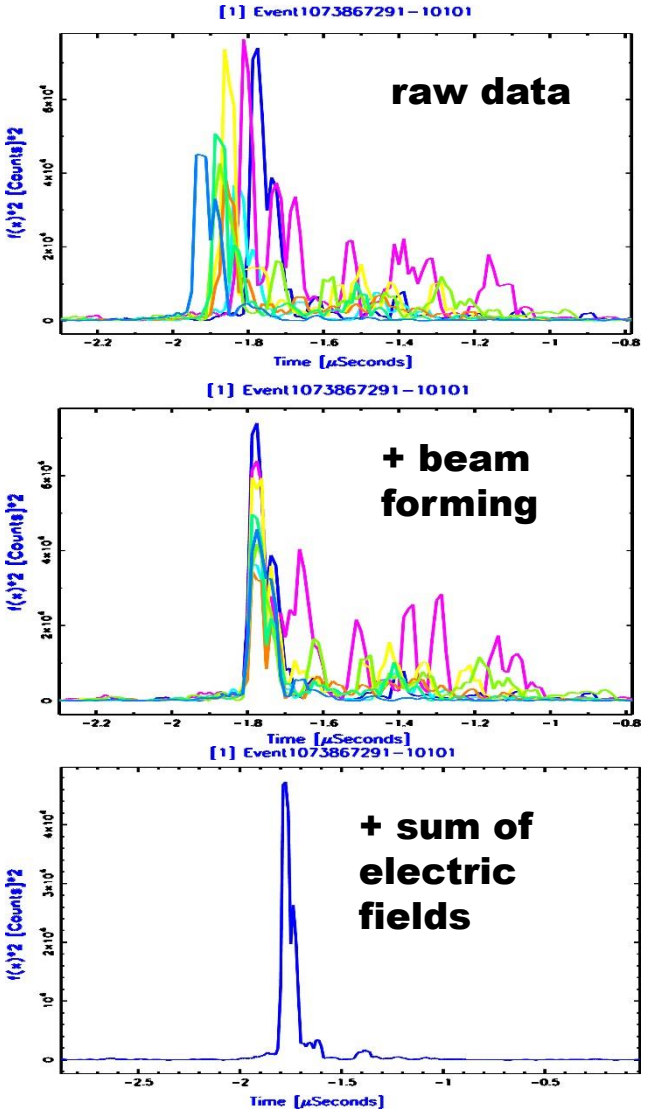
➔Development of a new detection technique!

LOPES: Proof of principle

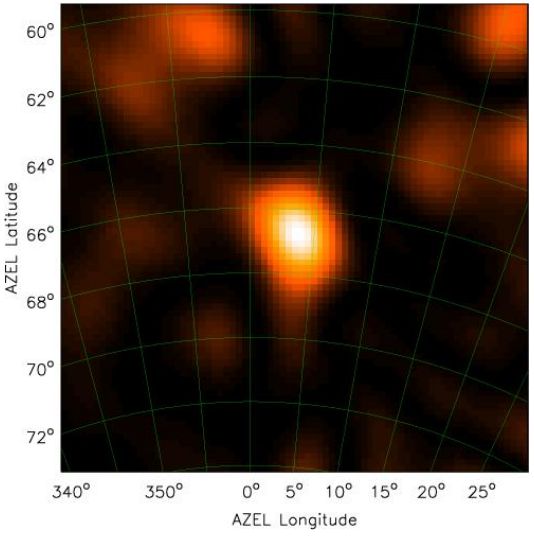
1. KASCADE measurement



2. Radio data analysis

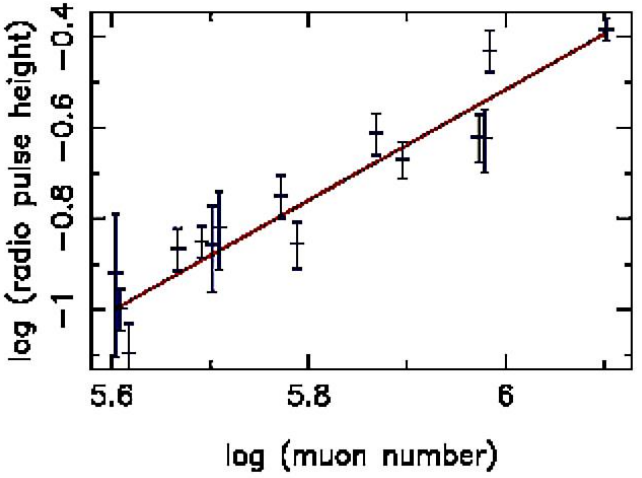


3. Skymapping



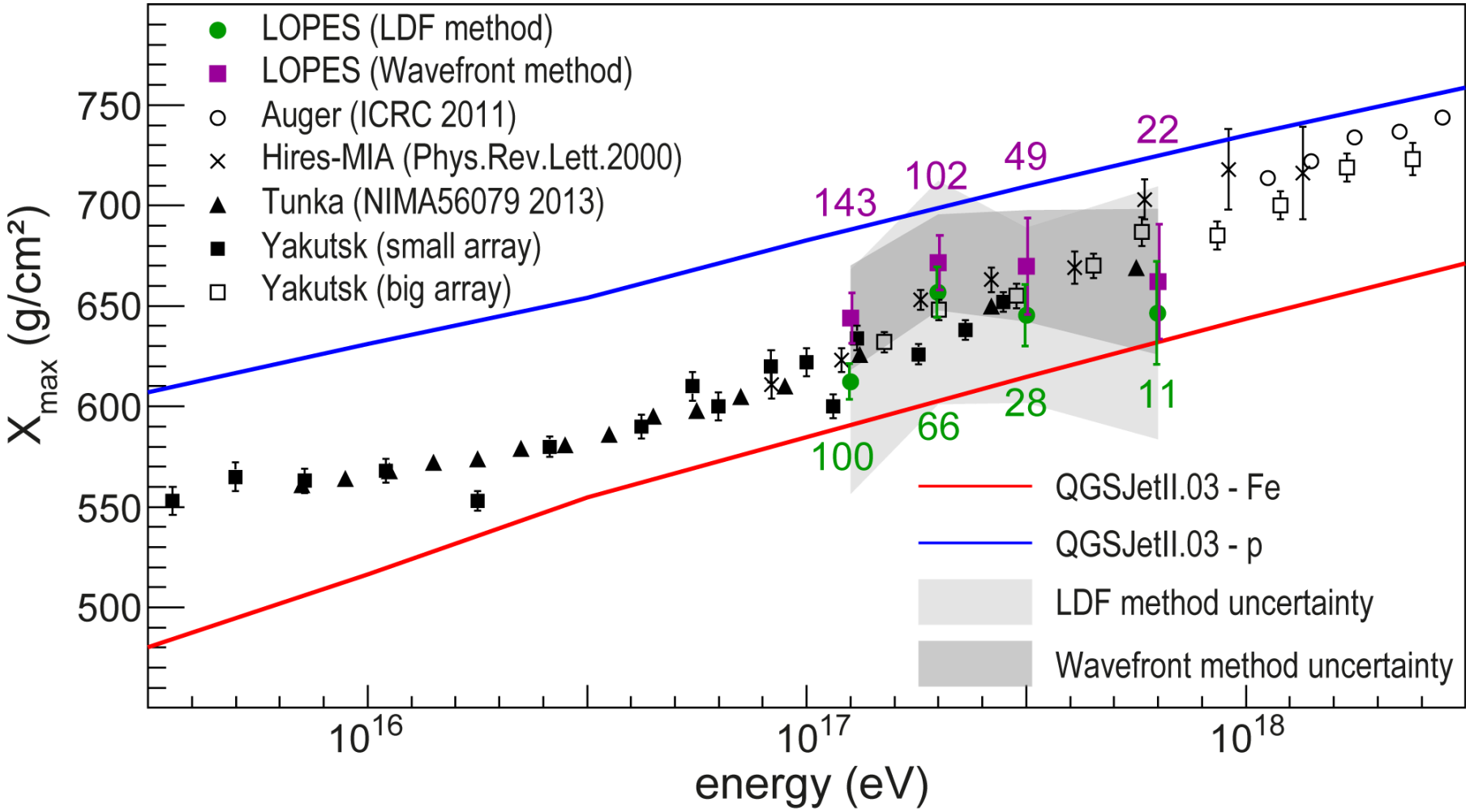
4. Many events

meanwhile >500 events



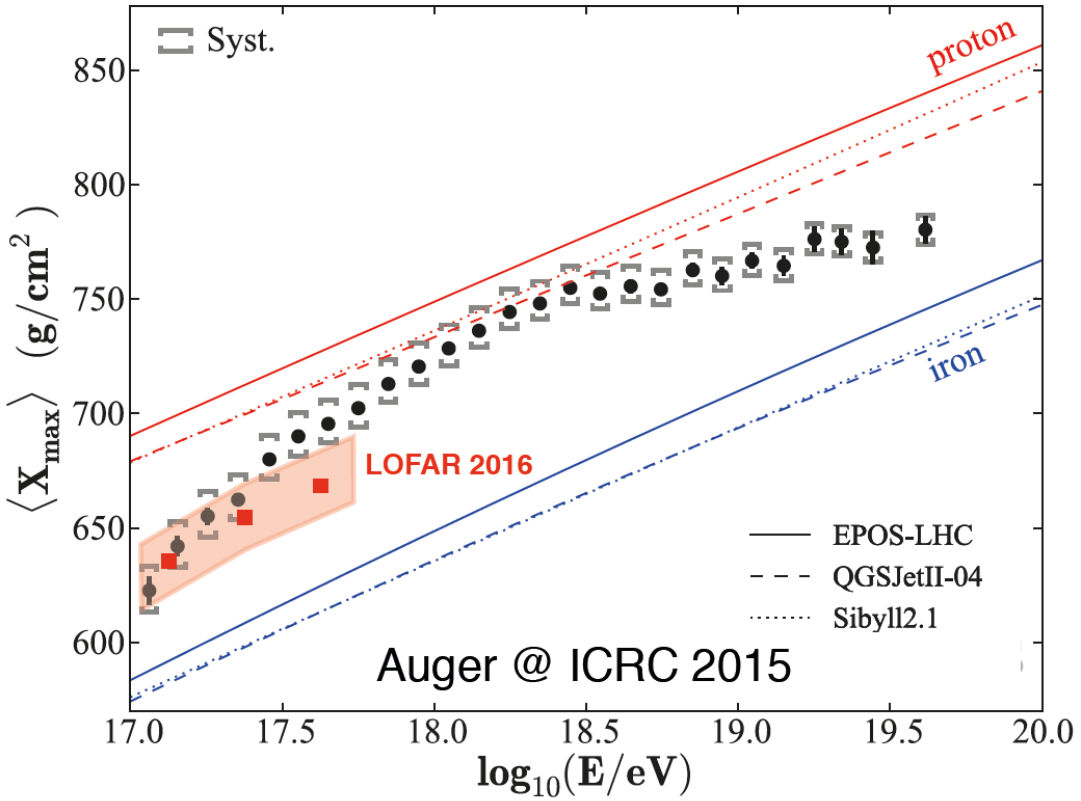
LOPES collaboration, Nature 425 (2005) 313

Composition measurements by LOPES



Xmax / Composition by Radio

A lot of (promising) progress in Xmax determination by radio Experiments



- **published already by**

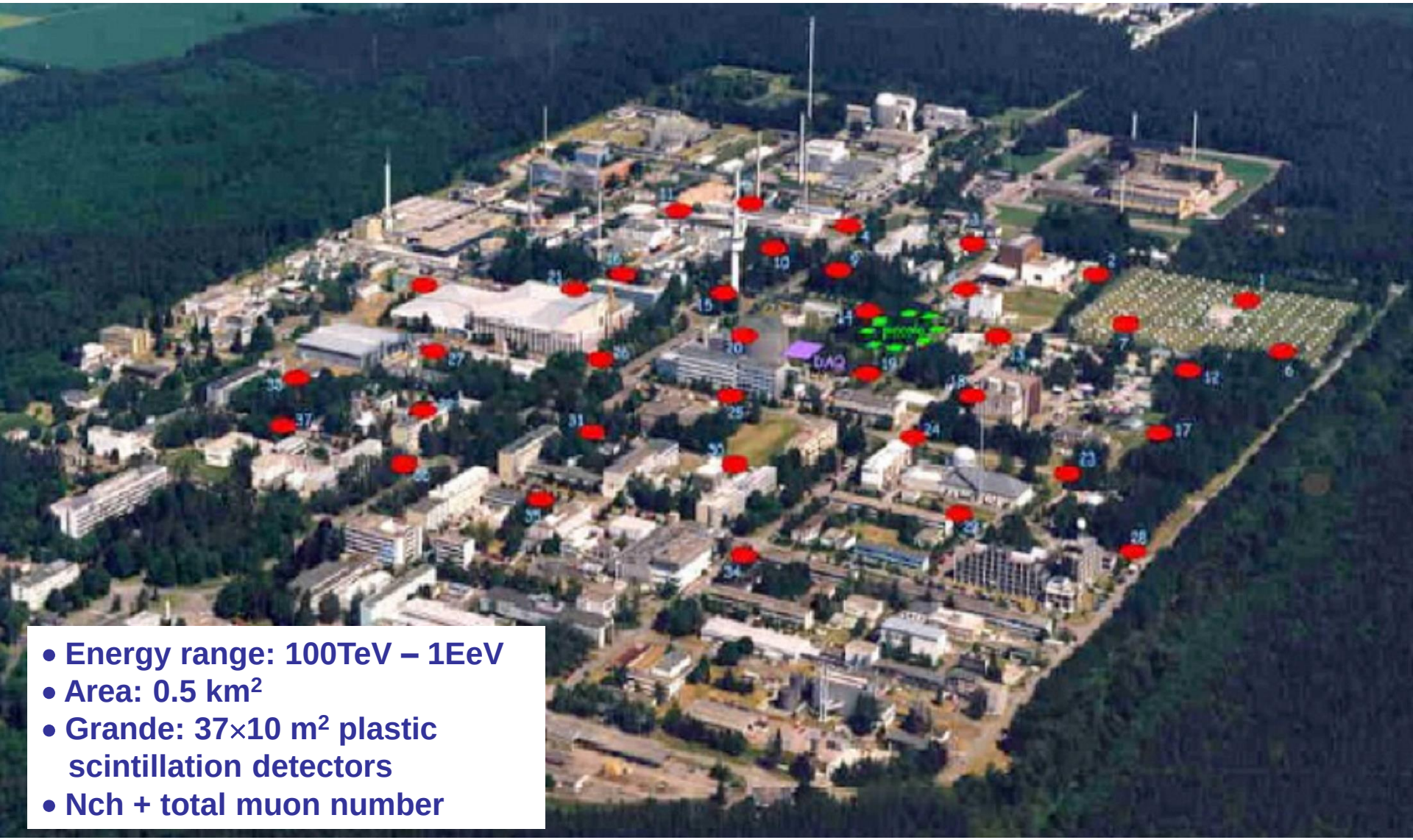
LOPES
PhysRevD 90(2014)062001
Tunka-Rex
JCAP 01(2016)052
LOFAR
Nature 531(2016)70

- **Auger/AERA promising**

- **Higher energy**
- **More accurate EAS**
- **Calibration**
- **Various methods**

➔ **Interpretation debatable:** “Unless, contrary to current expectations, the extragalactic component of cosmic rays contributes substantially to the total flux below $10^{17.5}$ eV, our measurements indicate the existence of an additional galactic component to account for the light composition we measured…..” (LOFAR@Nature)

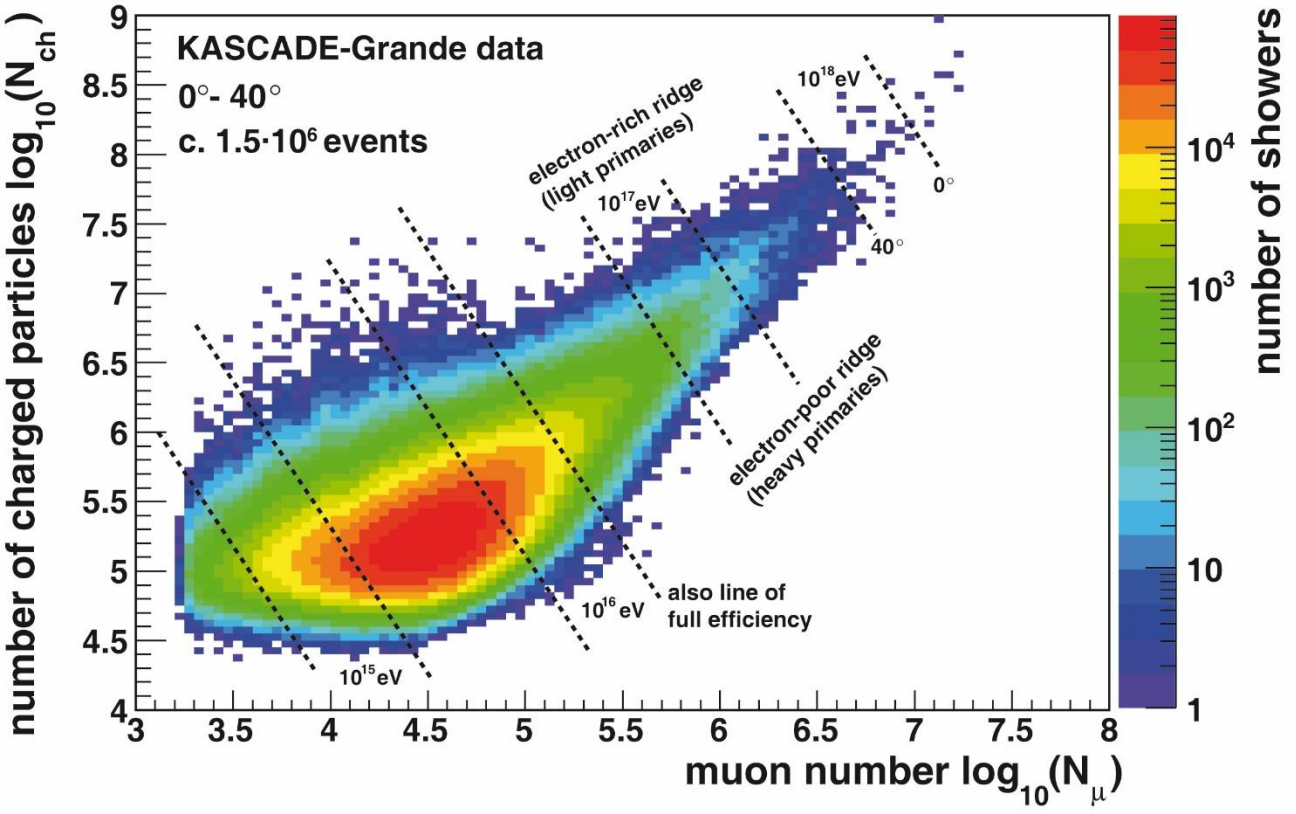
KASCADE-Grande



- Energy range: 100TeV – 1EeV
- Area: 0.5 km^2
- Grande: 37×10 m^2 plastic scintillation detectors
- Nch + total muon number

W.D.Apel et al, Nucl.Instr. and Meth. A620 (2010) 202

2-dimensional shower size spectrum



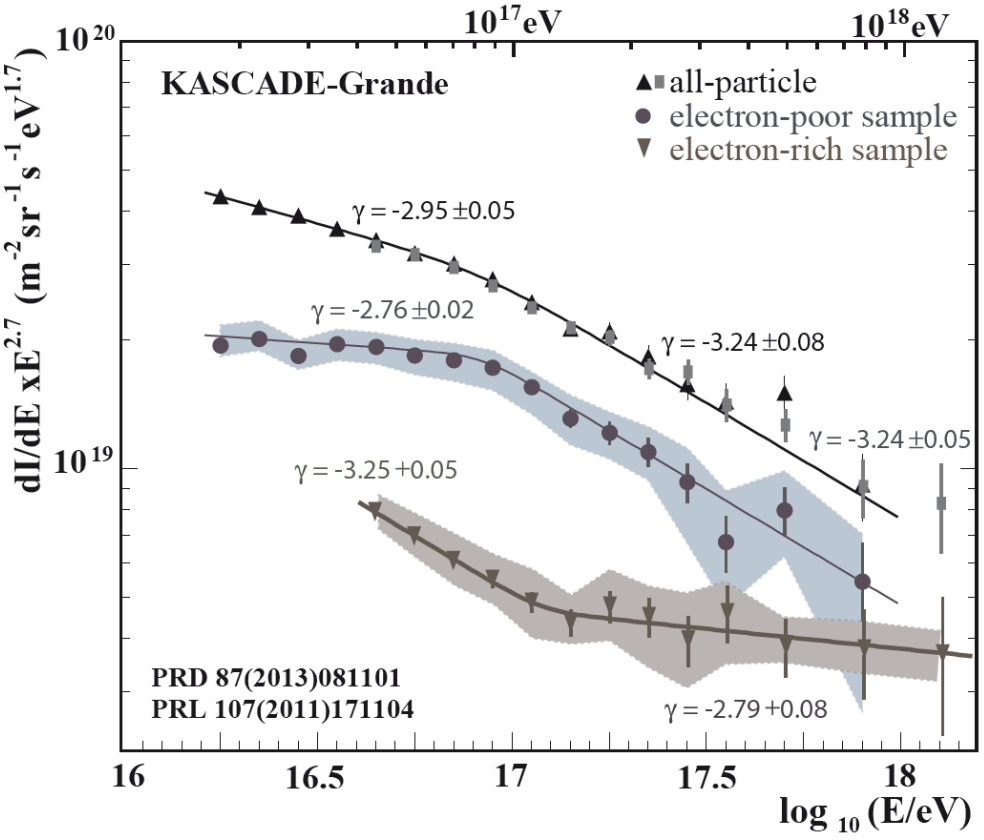
➔ **determination of primary energy**

➔ **separation in "electron-rich" and "electron-poor" event**

$$\log_{10}(E) = [a_p + (a_{Fe}-a_p)\cdot k] \cdot \log_{10}(N_{ch}) + b_p + (b_{Fe}-b_p)\cdot k$$

$$k = (\log_{10}(N_{ch}/N_\mu) - \log10(N_{ch}/N_\mu)_p) / (\log10(N_{ch}/N_\mu)_{Fe} - \log10(N_{ch}/N_\mu)_p)$$

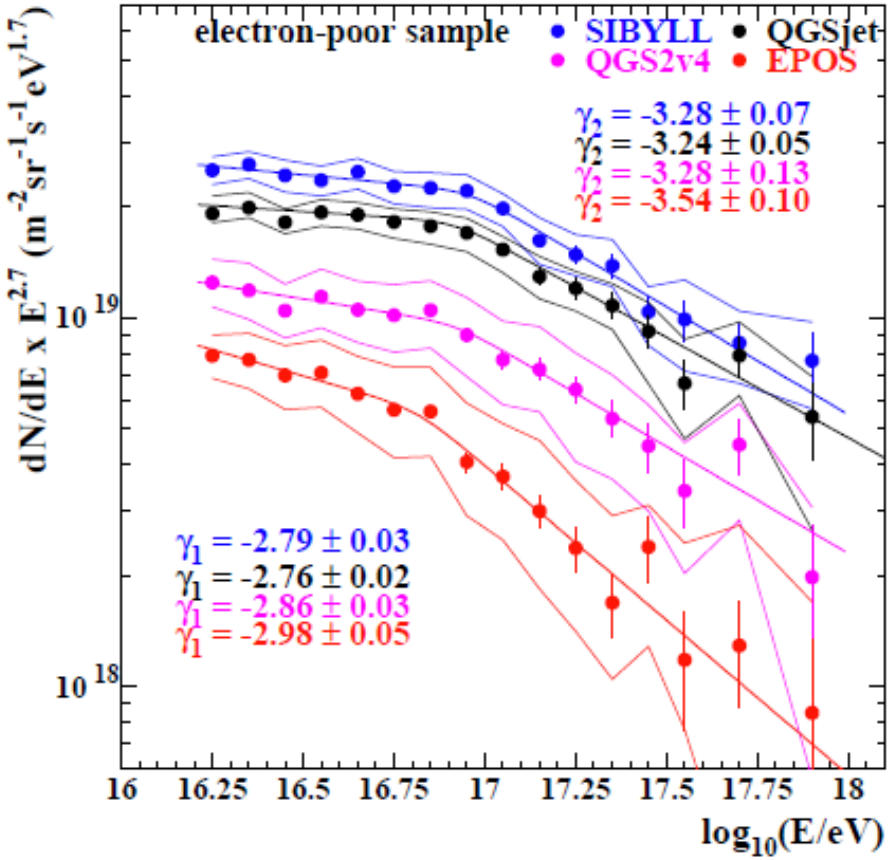
KASCADE-Grande
energy spectra of mass groups

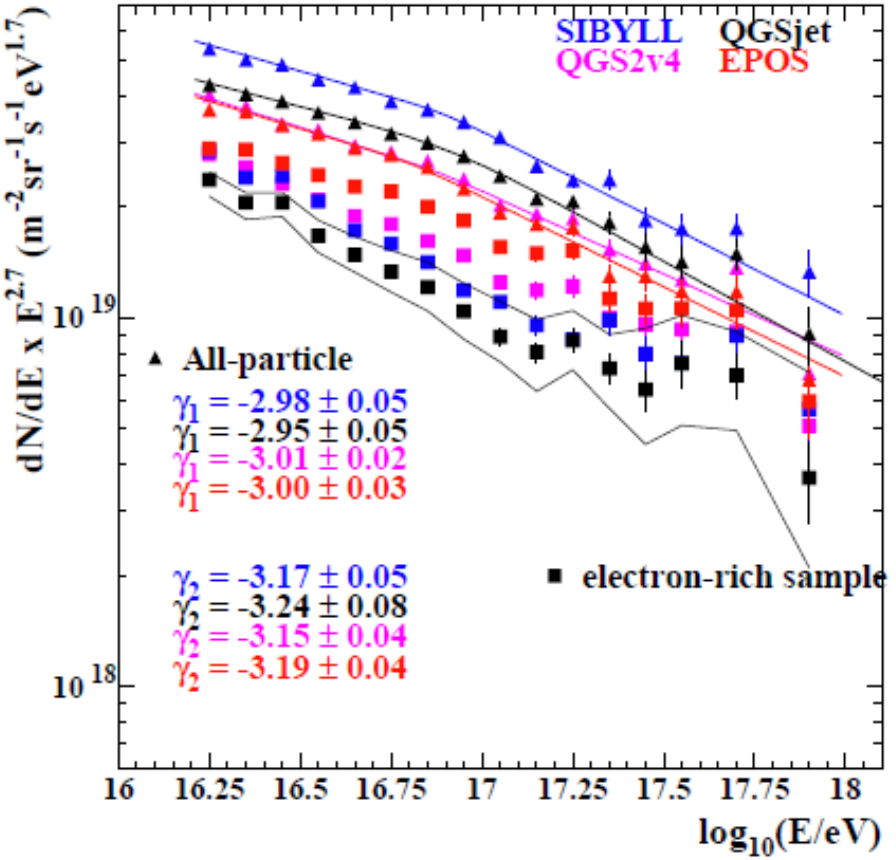


- **steepening due to heavy primaries (3.5σ)**
- **hardening at $10^{17.08}$ eV (5.8σ) in light spectrum**
- **slope change from γ = -3.25 to γ = -2.79!**

Phys.Rev.Lett. 107 (2011) 171104
Phys.Rev.D (R) 87 (2013) 081101

KASCADE-Grande: model dependence



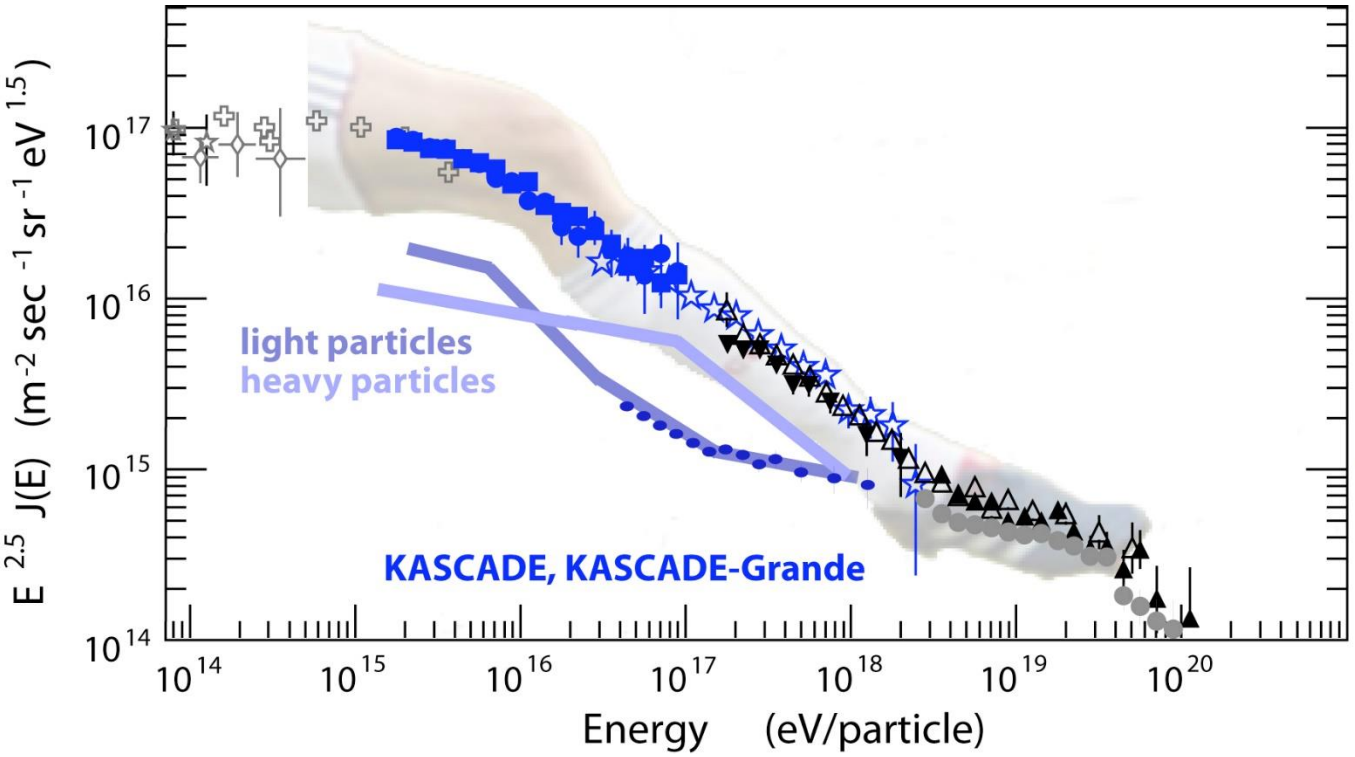


- **Structures of all-particle, heavy and light spectra similar**
  ➔ **knee by light component and heavy component; ankle by light component**
- **relative abundances different** **for different high-energy hadronic interaction models**

Advances in Space Research 53 (2014) 1456

30 March 2009 – official closure ceremony

Light and Heavy Knees, Ankles, and Transition



- ➔ **KASCADE: knee of light primaries at ~3·10^{15}eV**
- ➔ **Hardening at 10^{16}eV due to knee of medium component**
- ➔ **KASCADE-Grande: knee of heavy primaries at ~9·10^{16} eV**
- ➔ **heavy knee less distinct compared to light knee**
- ➔ **mixed composition for 10^{15} to ~ 8·10^{17} eV**
- ➔ **light ankle at 1-2·10^{17} eV**

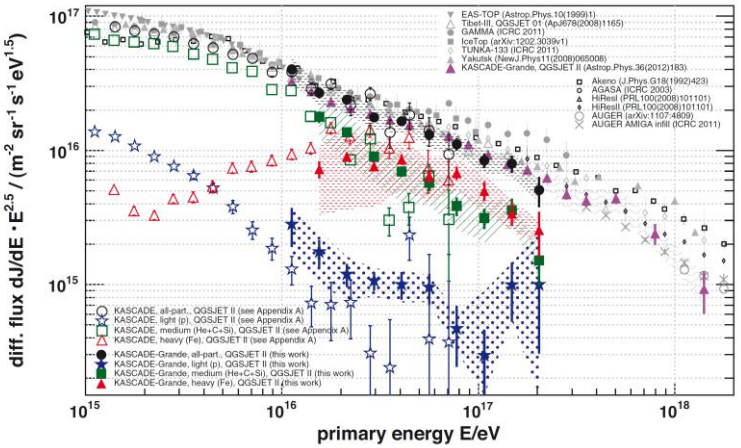
knee positions $\propto Z$

KASCADE-Grande: Next

- **KASCADE + KASCADE-Grande**
  - finally closed end 2012
  - now fully dismantled



- **combined analysis**
  - for coherent spectrum and composition 10^{14}-10^{18} eV

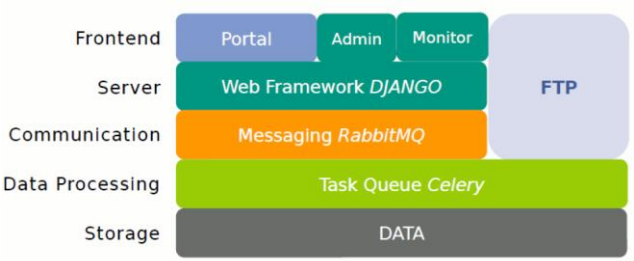


- **detailed data analysis** (20y high-quality data)
  - testing hadronic interaction models
  - anisotropy studies
  - radio (LOPES and CROME)

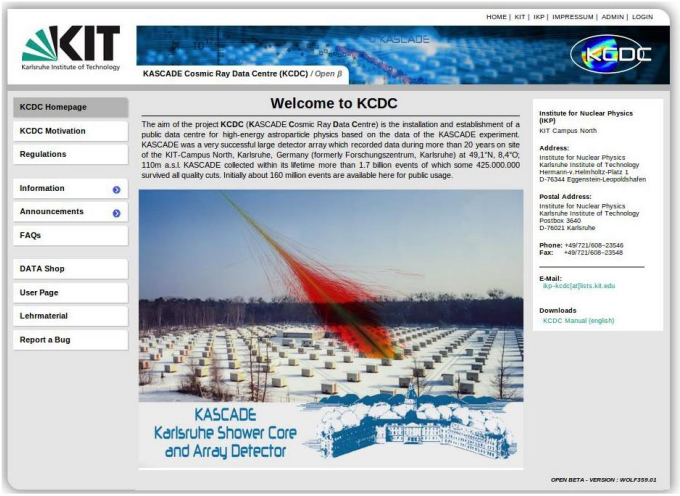
- **KCDC**
  - KASCADE Cosmic ray Data Centre

https://kcdc.ikp.kit.edu/

- KCDC = publishing research data
  from the KASCADE experiment
- Motivation and Idea of Open Data:
  general public has to be able to
  access and use the data
  the data has to be preserved for
  future generations



- Web portal:
  providing a modern software solution
  for publishing KASCADE data
  for a general audience
  In a second step: release the software
  as Open Source for free use by
  other experiments



- Data access:
  $1.6 \cdot 10^8$ EAS events of first data release
  is now available

Paper in preparation

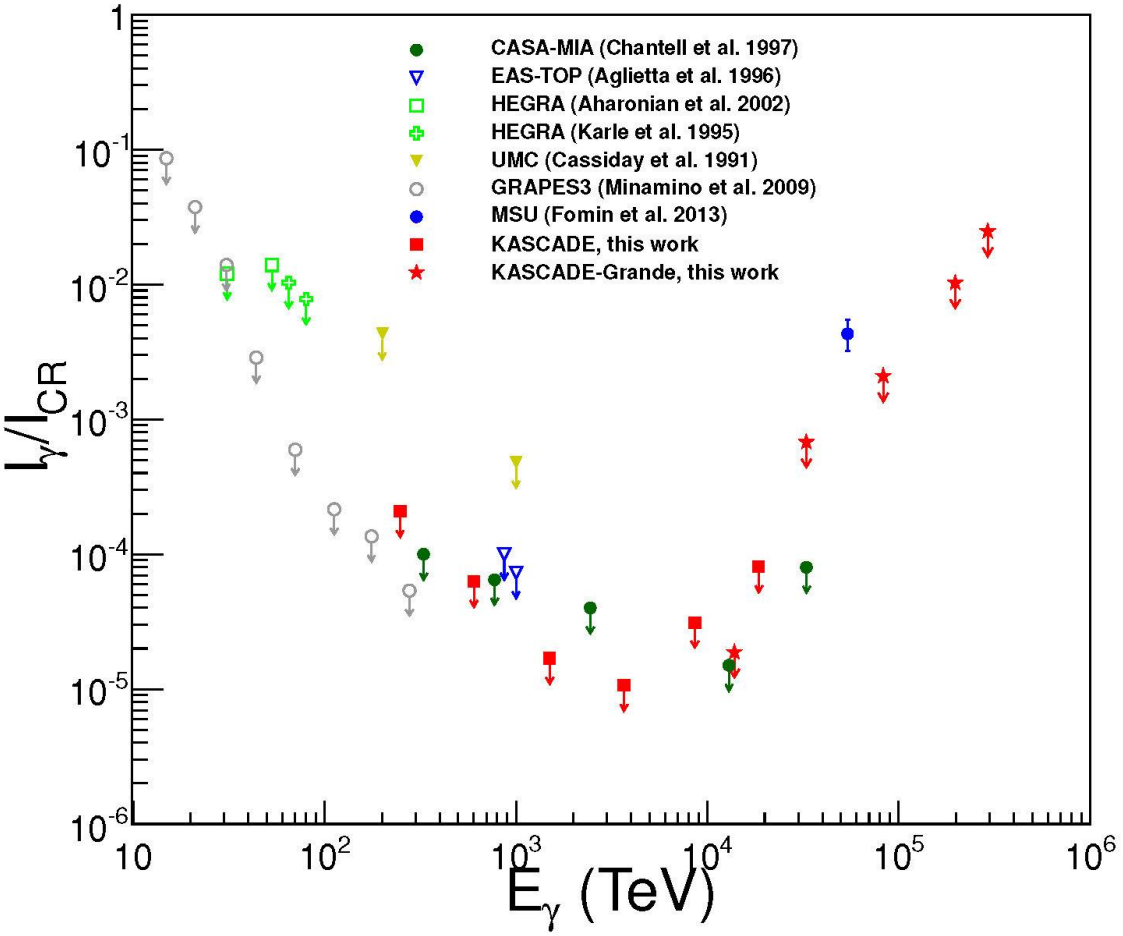
KASCADE-Grande: Mission Accomplished !!

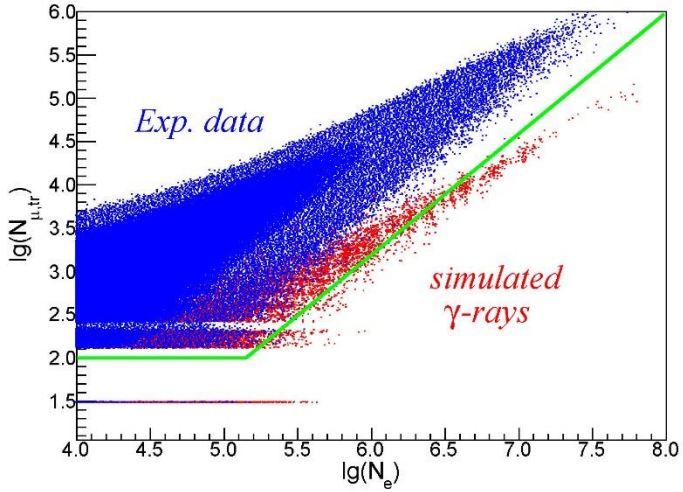
open access to research data
https://kcdc.ikp.kit.edu

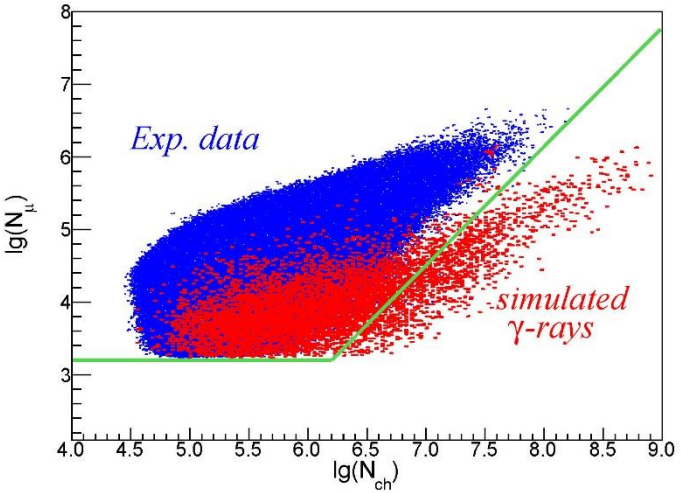
KASCADE-Grande: Mission Accomplished !!

open access to research data
https://kcdc.ikp.kit.edu

Ratio of diffuse Gamma-ray Flux to cosmic ray flux



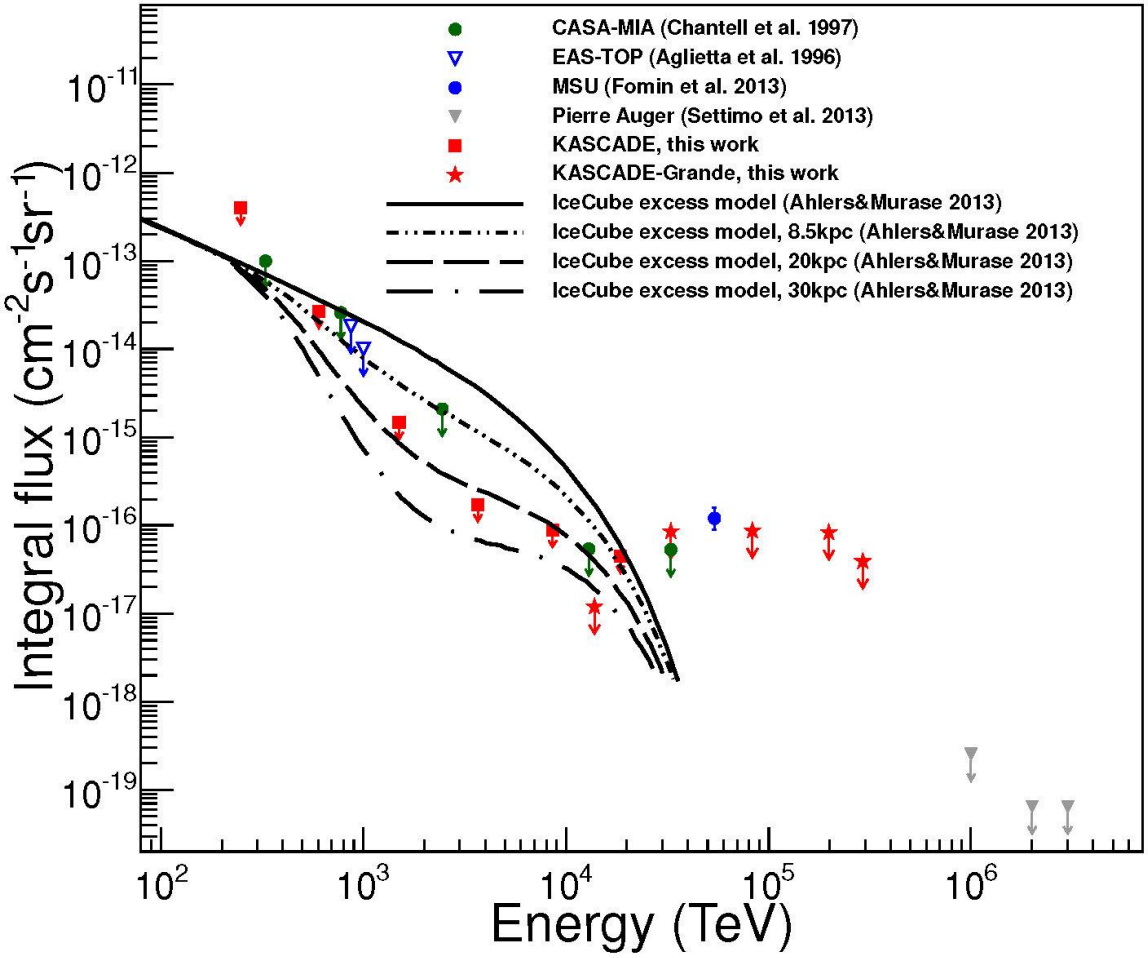




- **obtained by looking for muon-poor air shower**

Paper in preparation
Analysis by Donghwa Kang

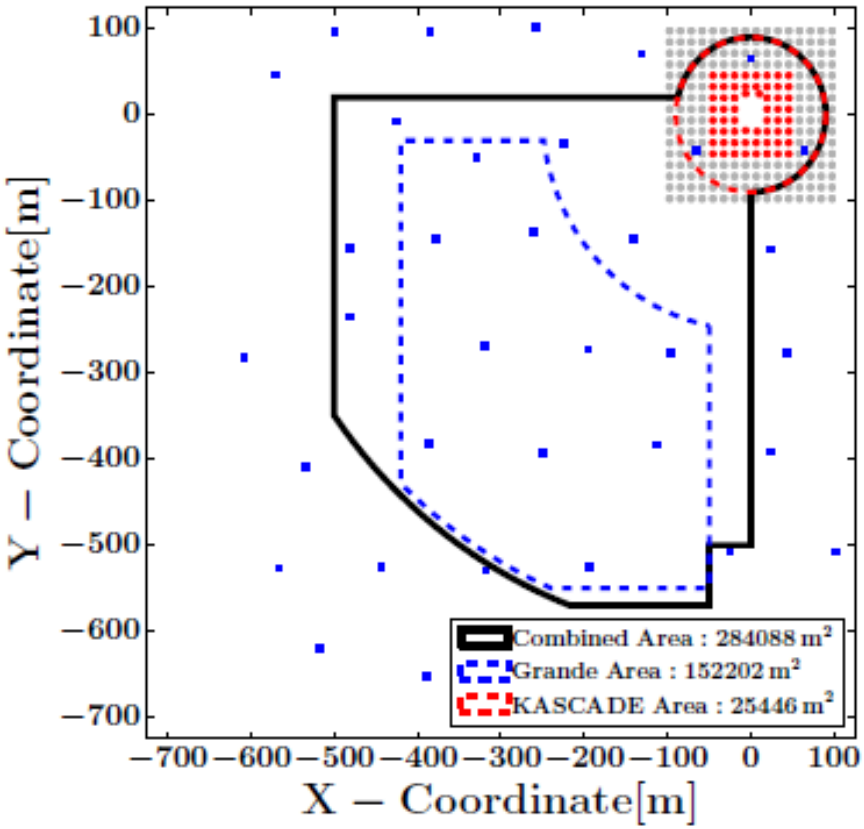
Limits on diffuse Gamma-ray Flux

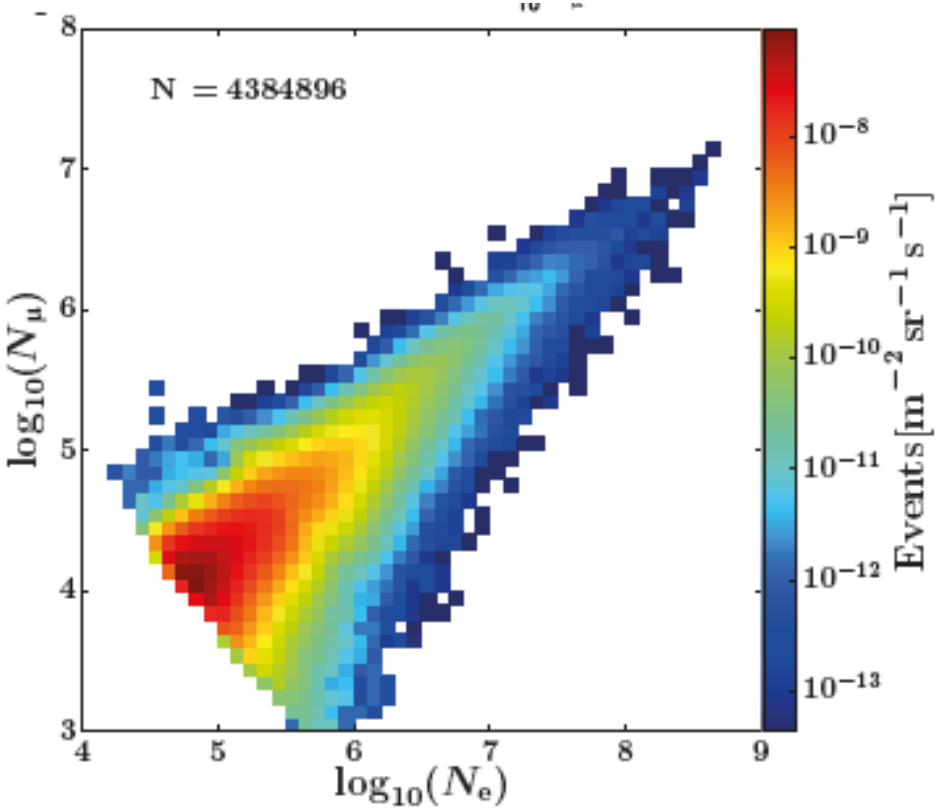


- **limits on diffuse Gamma-ray flux constrain the origin of IceCube-neutrinos**

← Reject the model of IceCube excess coming from <20kpc in the galaxy

KASCADE-Grande: Combined Analysis





- **for KASCADE: additional stations at larger distances**

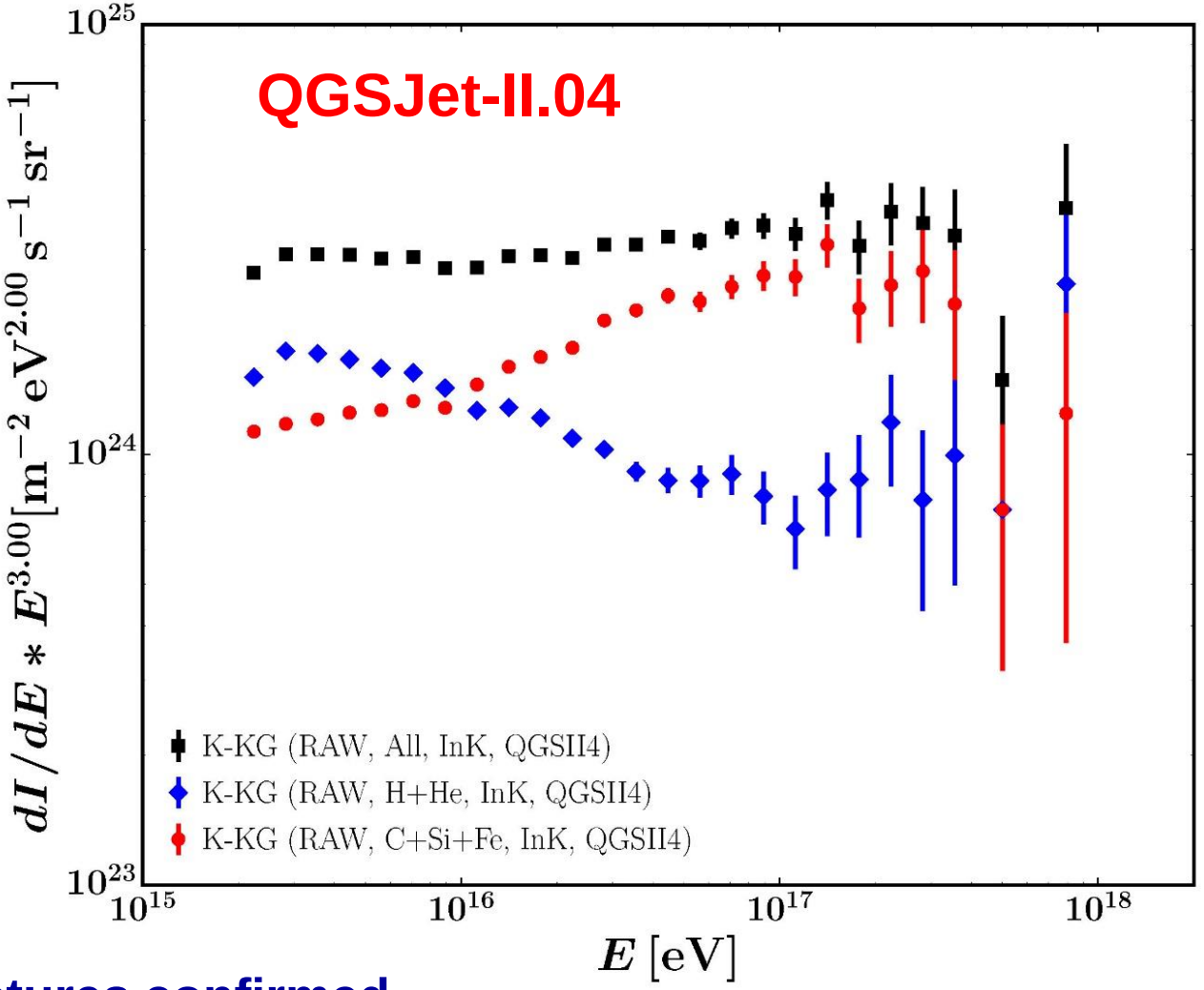
  ➔ higher energies

- **for Grande: additional 252 stations**

  ➔ higher accuracy

Analysis by Sven Schoo

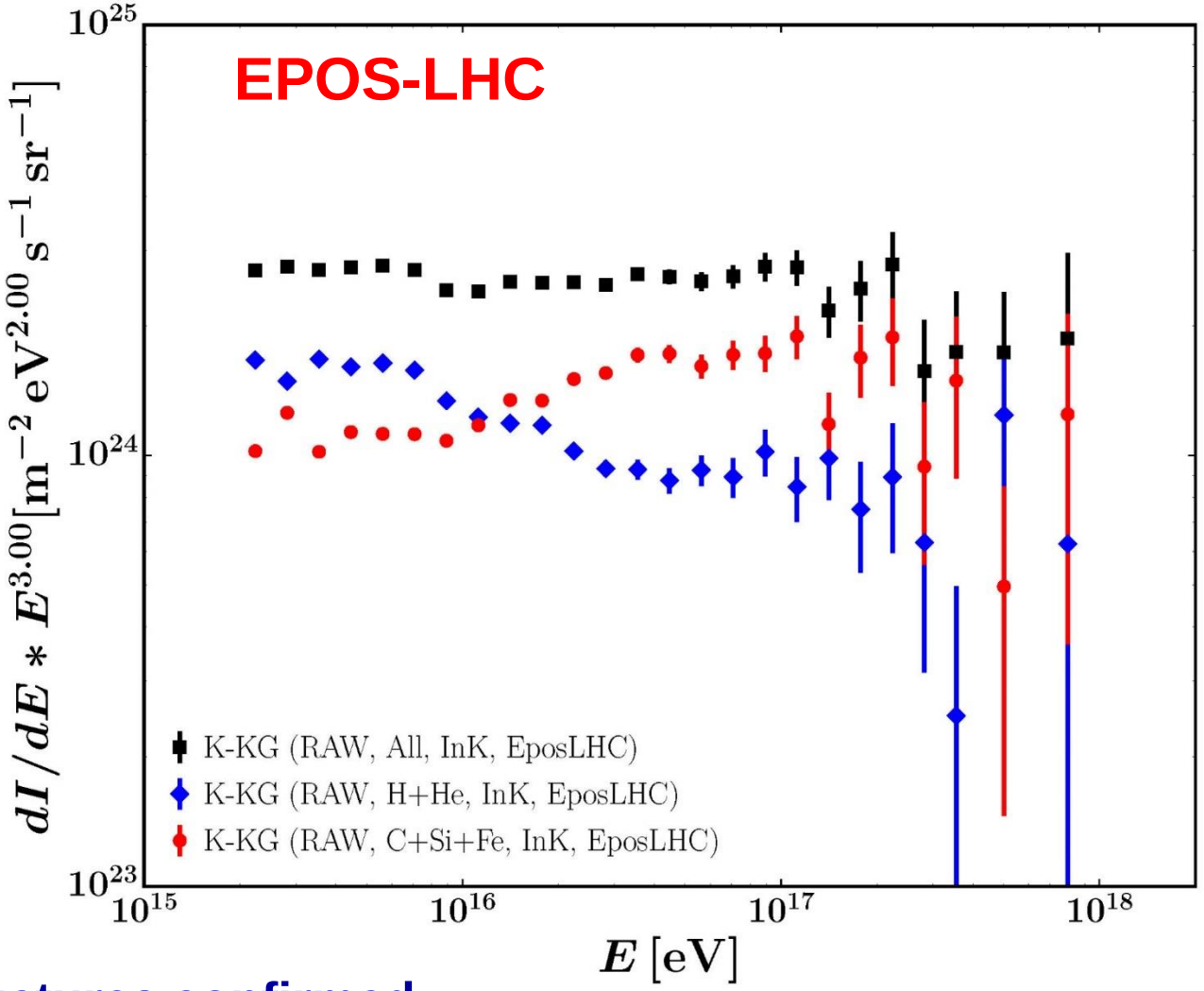
KASCADE-Grande: Combined Analysis resulting energy spectra



- **all structures confirmed**

Spectra not corrected for uncertainties

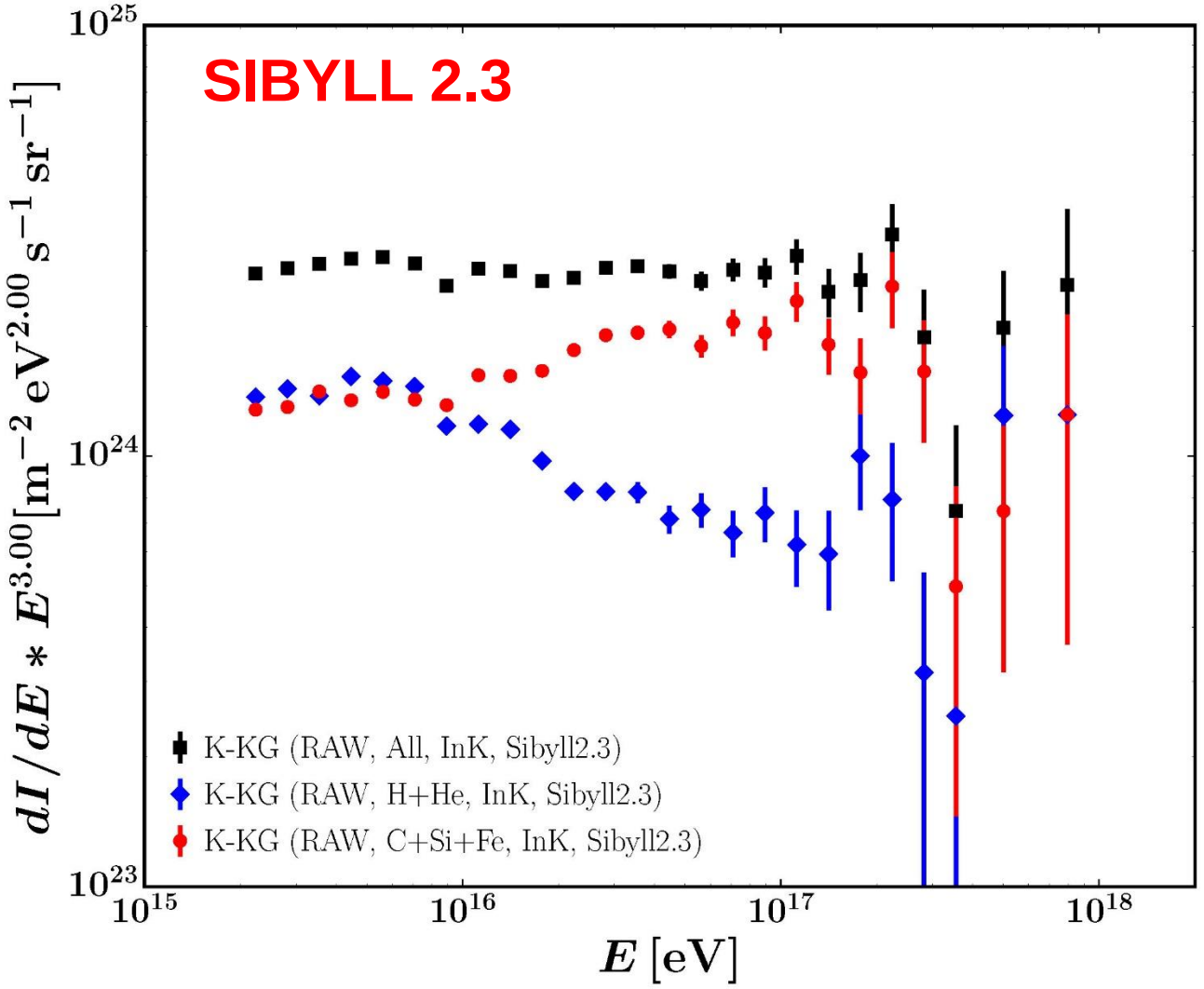
KASCADE-Grande: Combined Analysis resulting energy spectra



- **all structures confirmed**

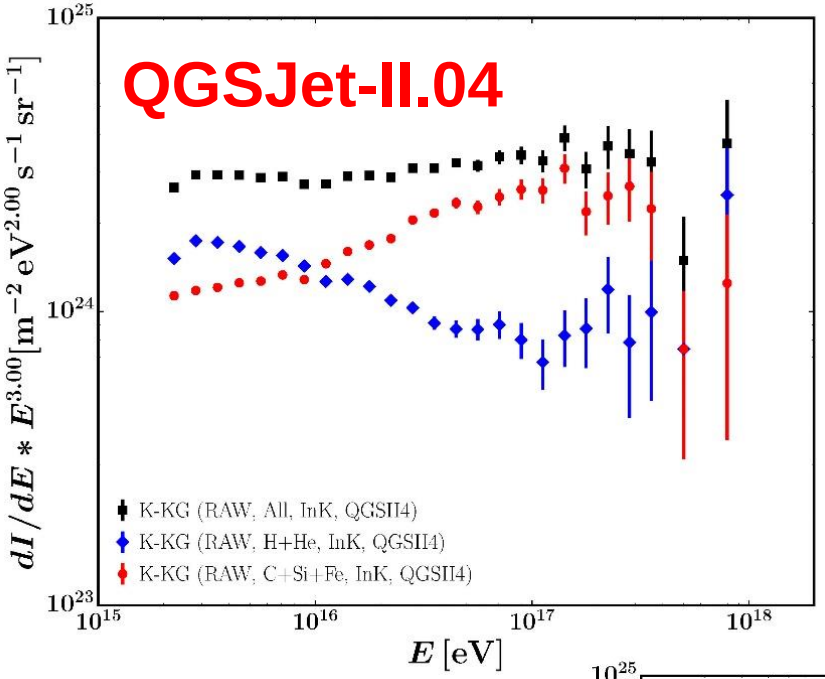
Spectra not corrected for uncertainties

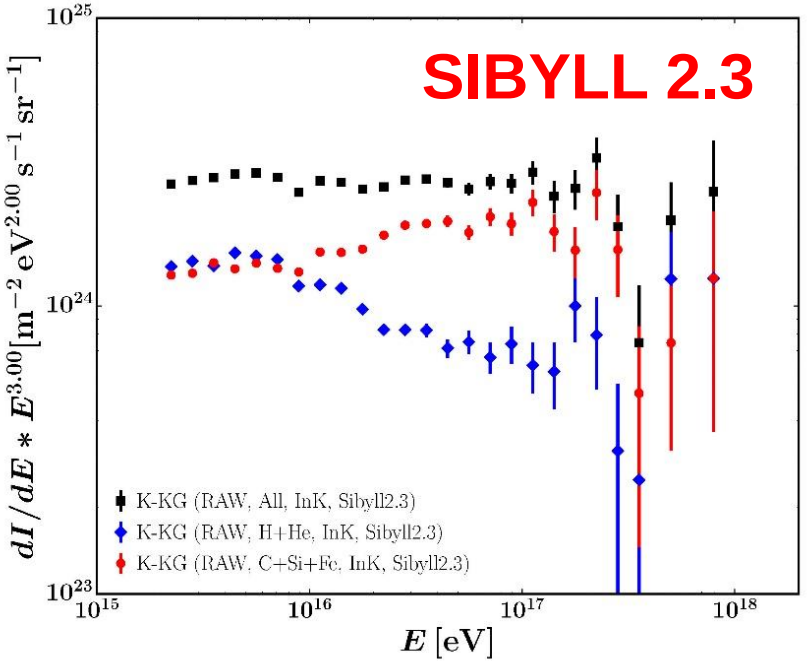
KASCADE-Grande: Combined Analysis resulting energy spectra

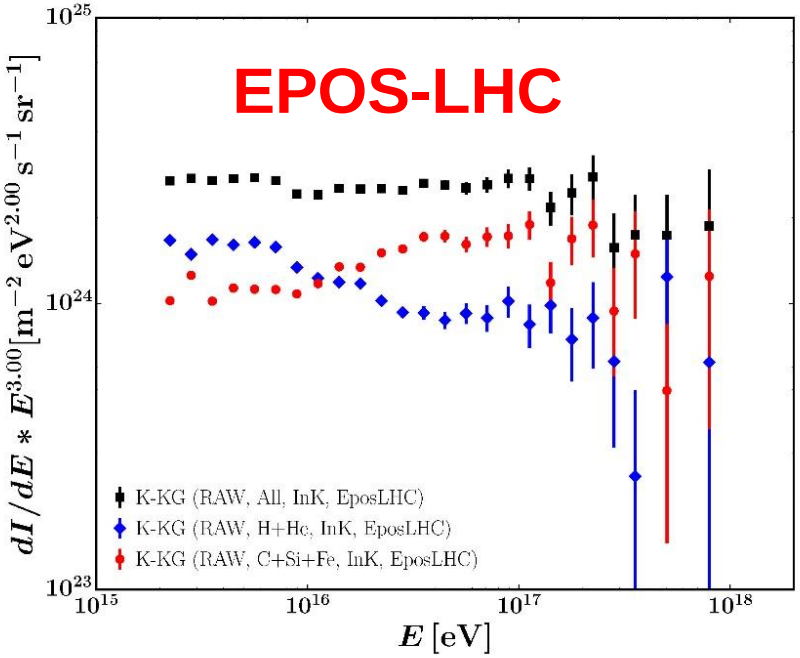


- **all structures confirmed**

Spectra not corrected for uncertainties



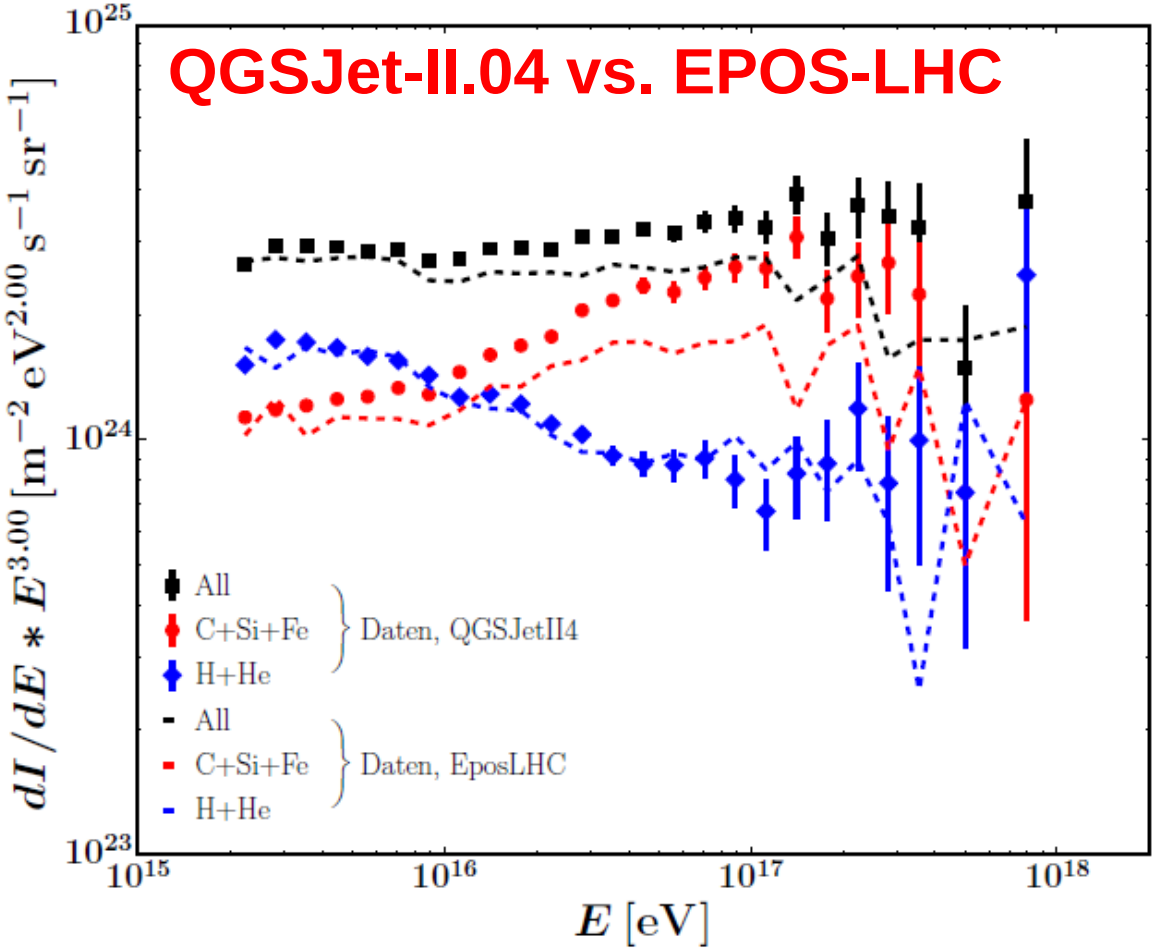




- **structures confirmed**
- **all particle spectrum good agreement**
- **relative abundance of light and heavy quite different**

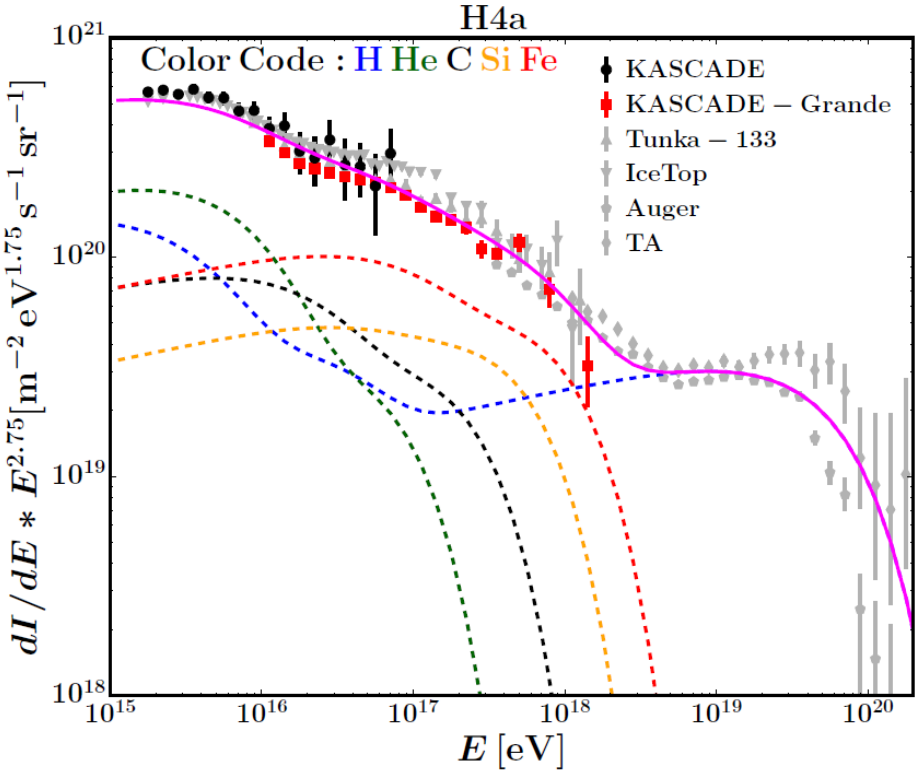
KASCADE-Grande: Combined Analysis

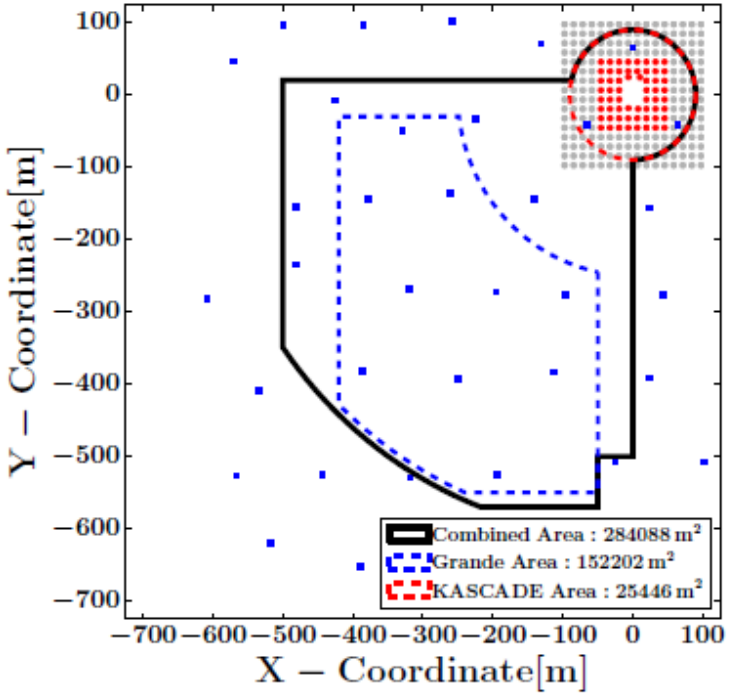
resulting energy spectra based on two hadronic interaction models



- **Post LHC models**
  - **light primary interactions okay?**
  - **heavy primary interactions show differences**

KASCADE-Grande: combined analysis
Check Hadronic Interaction Models





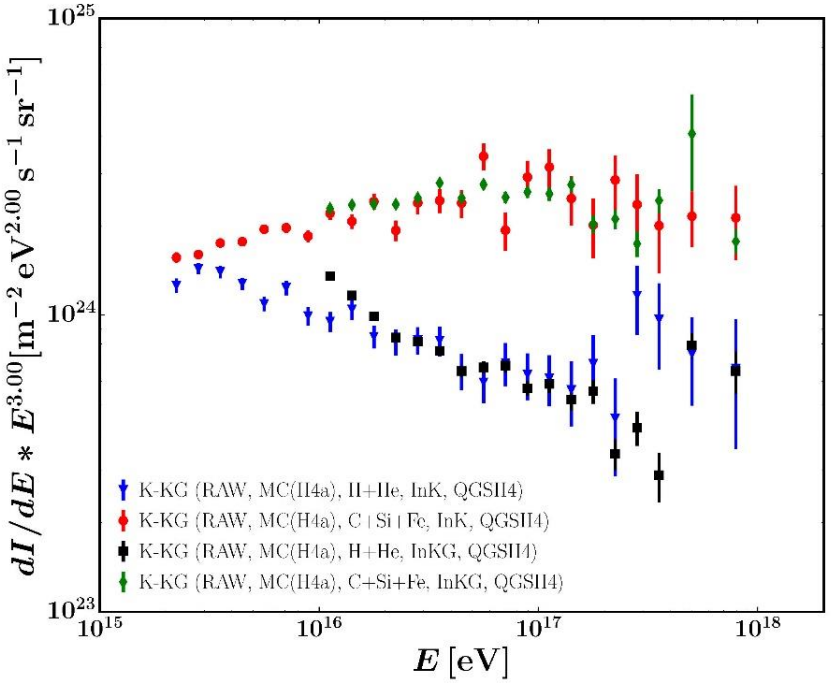
- **assume a composition model: H4a by Tom Gaisser**

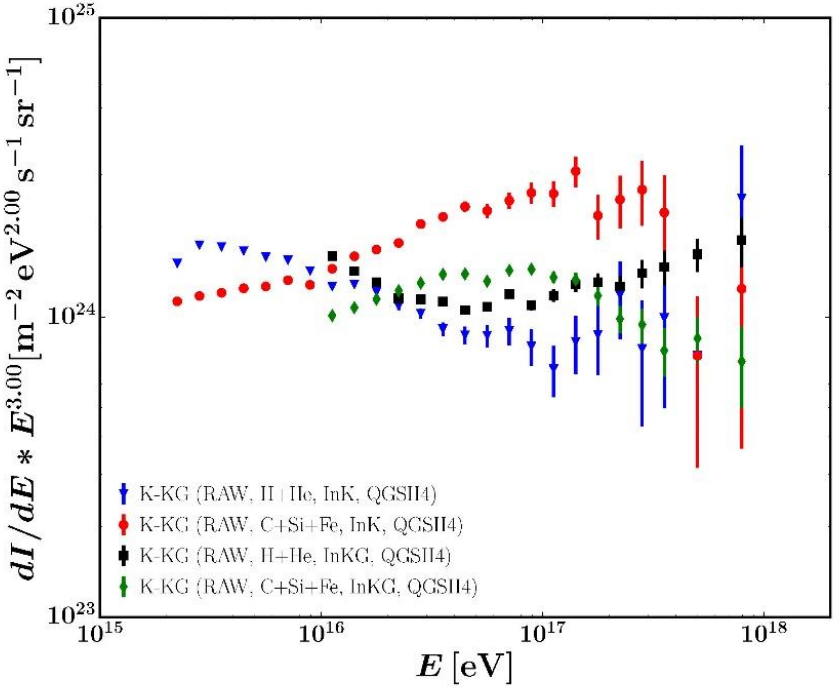
- **two selections: core located in KASCADE, core located in Grande**
  ➔ we measure “different” muons

KASCADE-Grande: Combined Analysis

Test of models

QGSJet-II.04





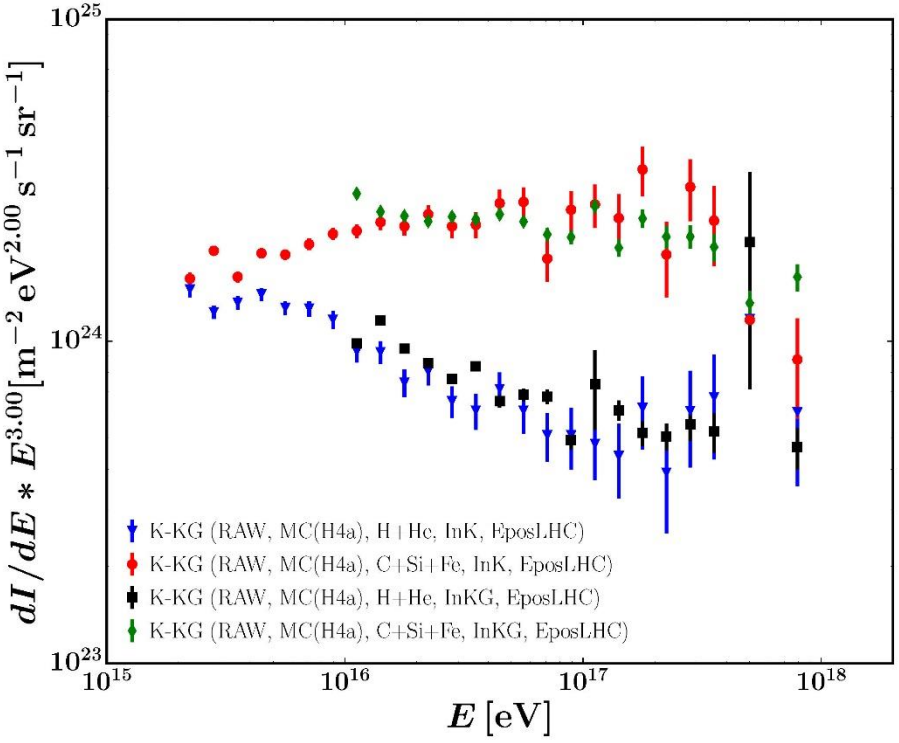
- **One model, but two selections:**
  Simulations okay, but for the data strong differences

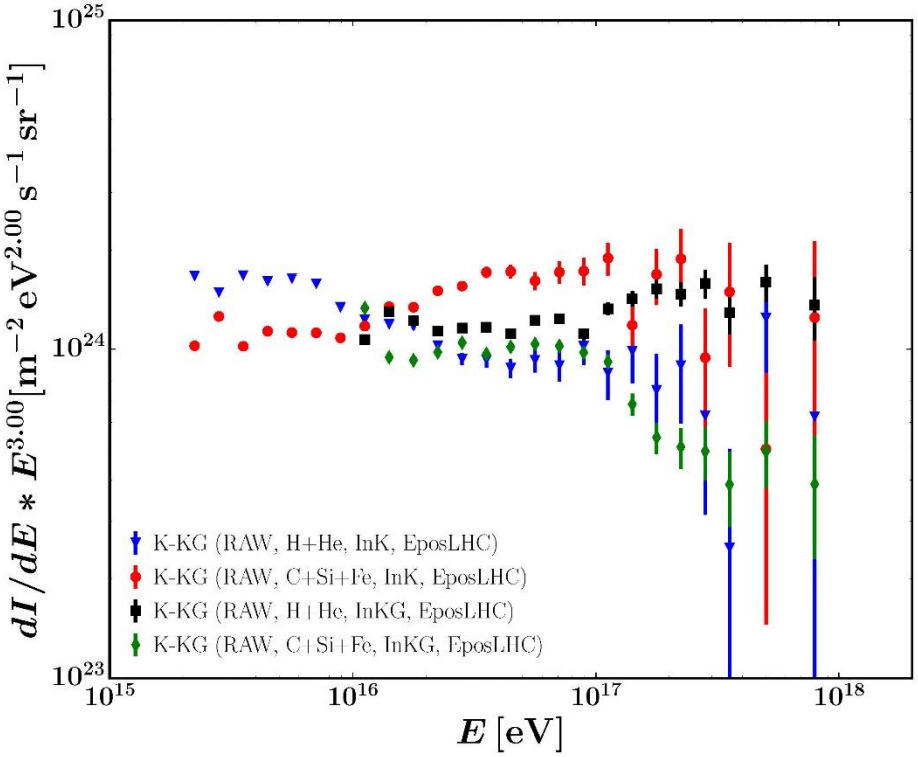
➔ **Muon component not sufficiently described**

KASCADE-Grande: Combined Analysis

Test of models

EPOS-LHC





- **One model, but two selections:**
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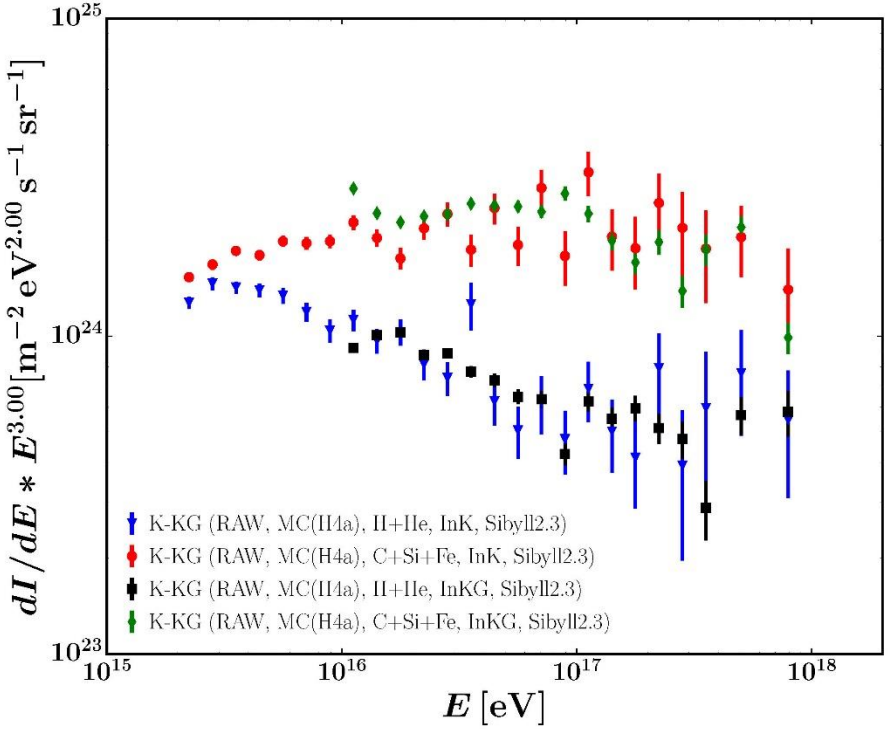
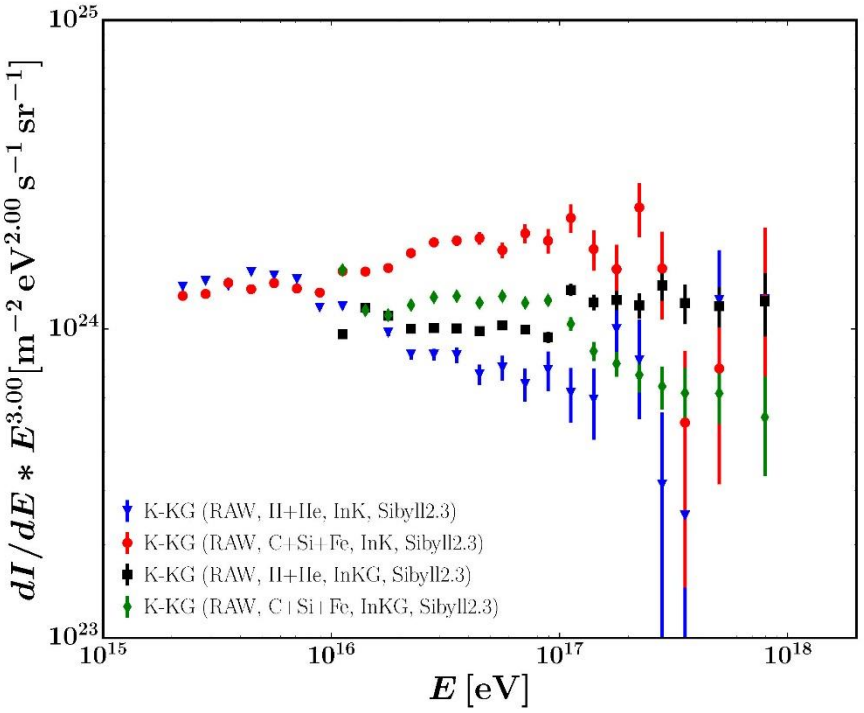
➔ **Muon component not sufficiently described**

KASCADE-Grande: Combined Analysis

Test of models

SIBYLL 2.3



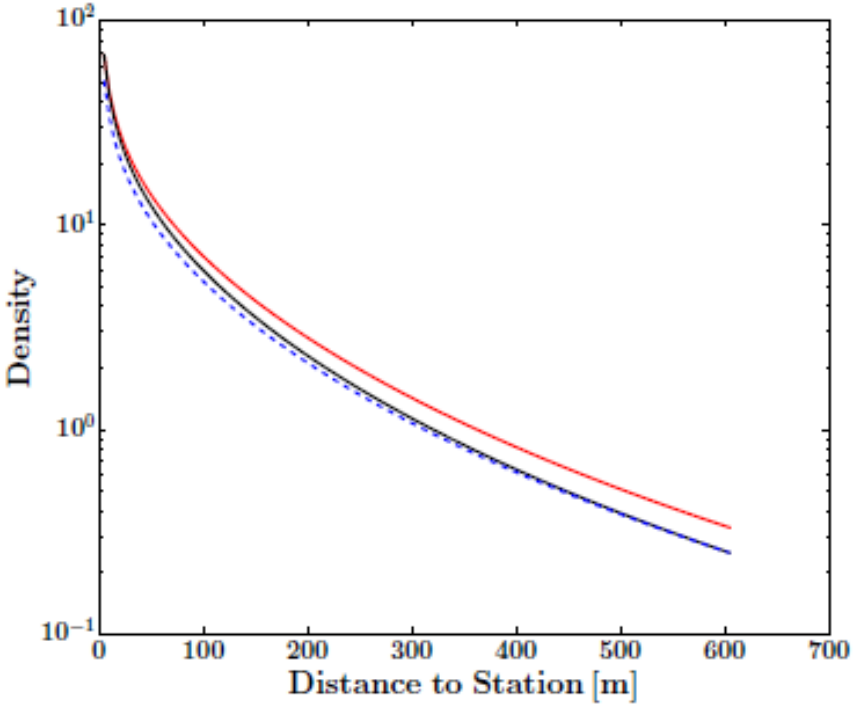


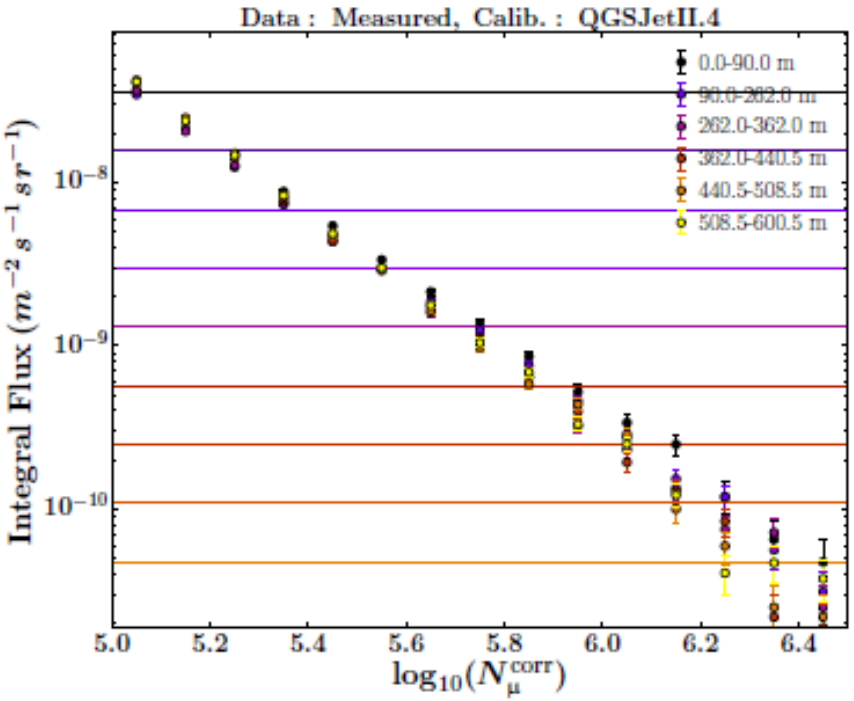
- **One model, but two selections:**
  Simulations okay, but for the data strong differences

➔ **Muon component not sufficiently described**

KASCADE-Grande: Combined Analysis

Test of models





Slope is parametrized with N_e, but fixed and too flat

Distance range covered by Muon detectors limited

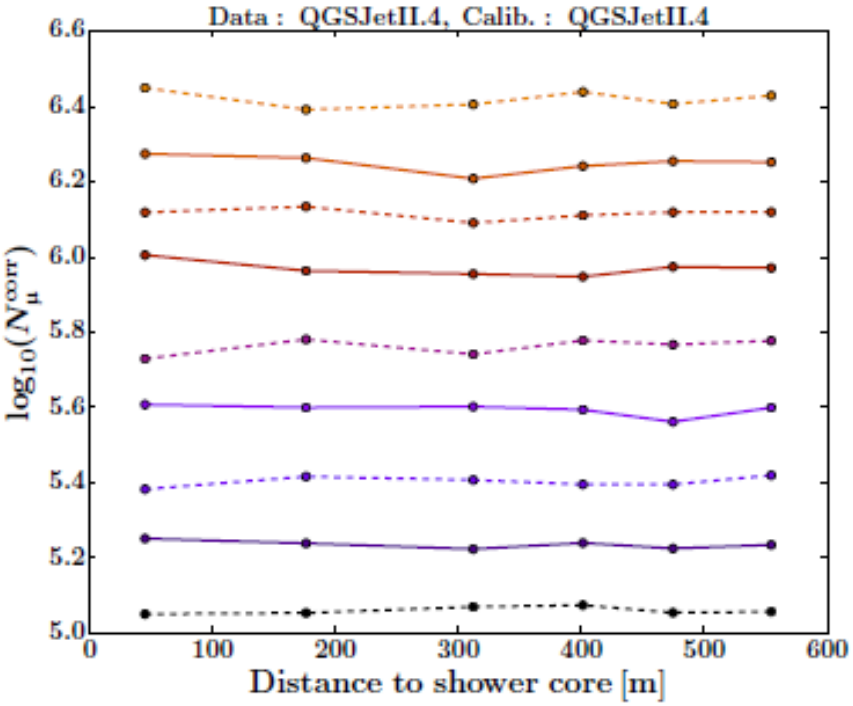
If core within KASCADE → too many muons.

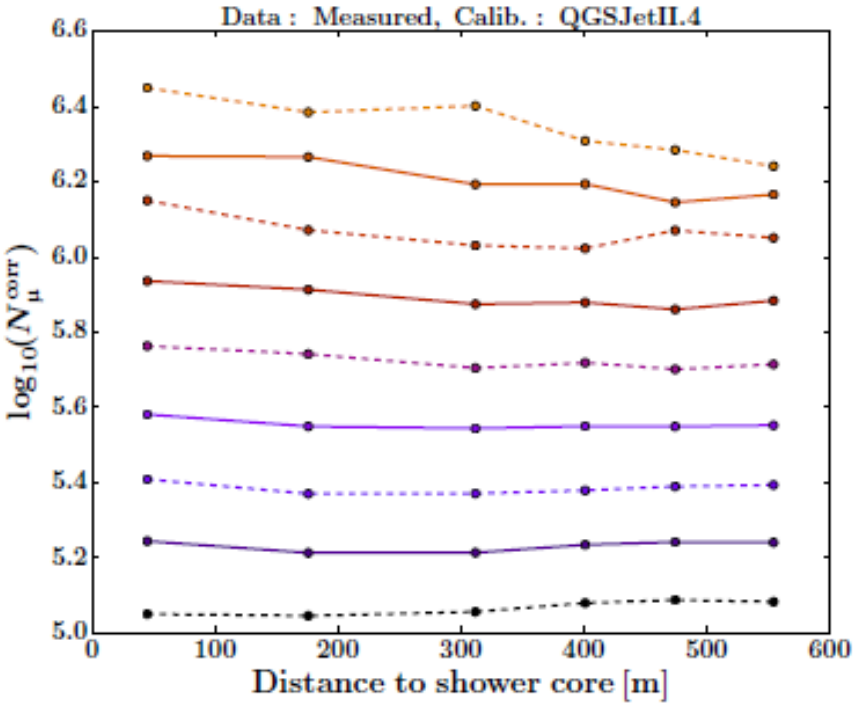
If core within Grande → too few muons.

Crosscheck using “Constant Intensity”

Sven Schoo

KASCADE-Grande: Combined Analysis
Test of models





Number of muons relatively constant for simulations

Number of muons corresponding to same intensity drops for measured data towards higher energies

Can be used to mitigate effect, however, more accurate MC needed

Sven Schoo

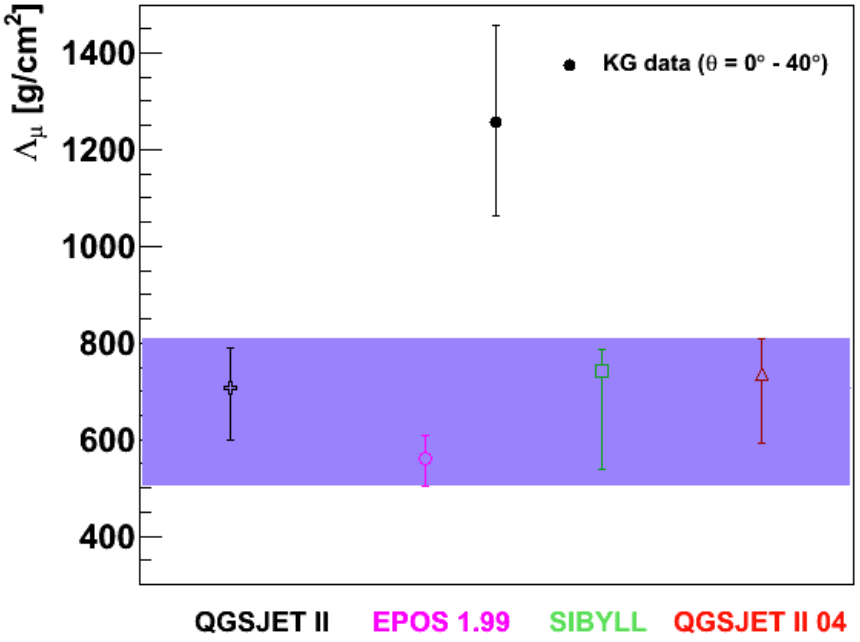
KASCADE-Grande: Muon Attenuation Length

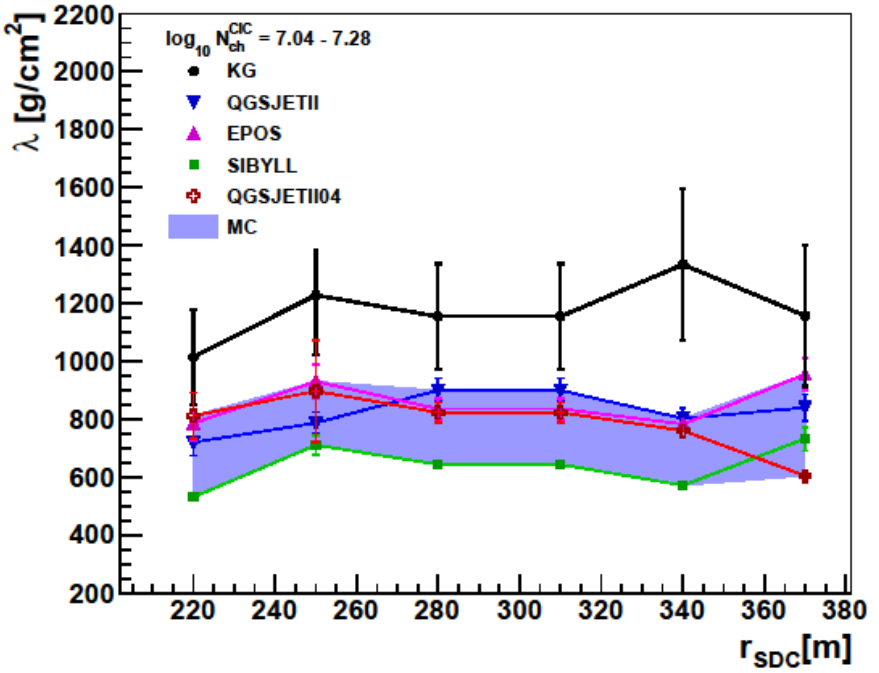
total muon number

$$N_\mu = N_{\mu,0} \exp[- X_0 \sec(\theta) / \Lambda_\mu]$$

local muon density

$$\rho_\mu(r) = \rho_{\mu,0}(r) \exp[- X_0 \sec(\theta) / \lambda_\mu(r)]$$





attenuation length measured is different from the predictions of Monte Carlo

➔ observed evolution of the muon content of EAS in the atmosphere is not described by the hadronic interaction models

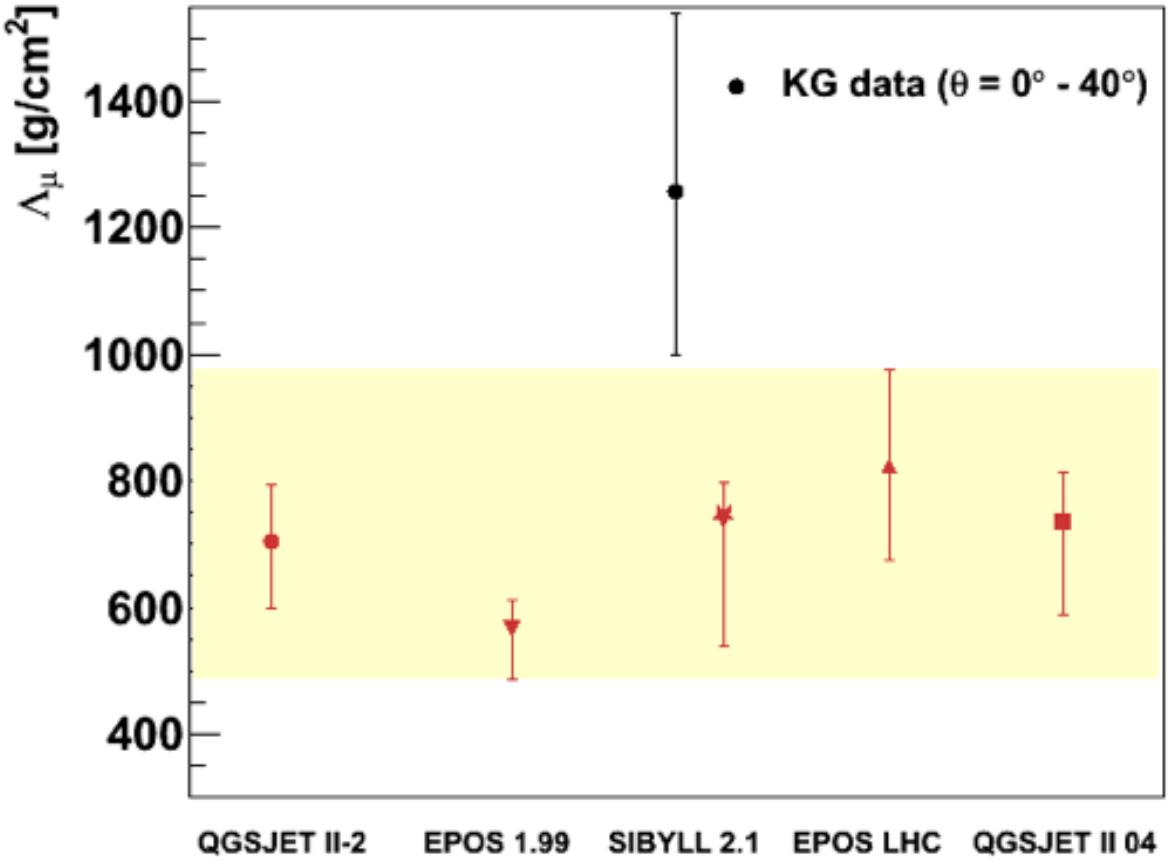
➔ influences absolute energy and mass scale, but not spectral features

KASCADE-Grande, ICRC 2013 #0772, paper in preparation

Juan Carlos Arteaga

KASCADE-Grande: Muon Attenuation Length

total muon number : $N_\mu = N_{\mu,0} \exp[- X_0 \sec(\theta) / \Lambda_\mu]$

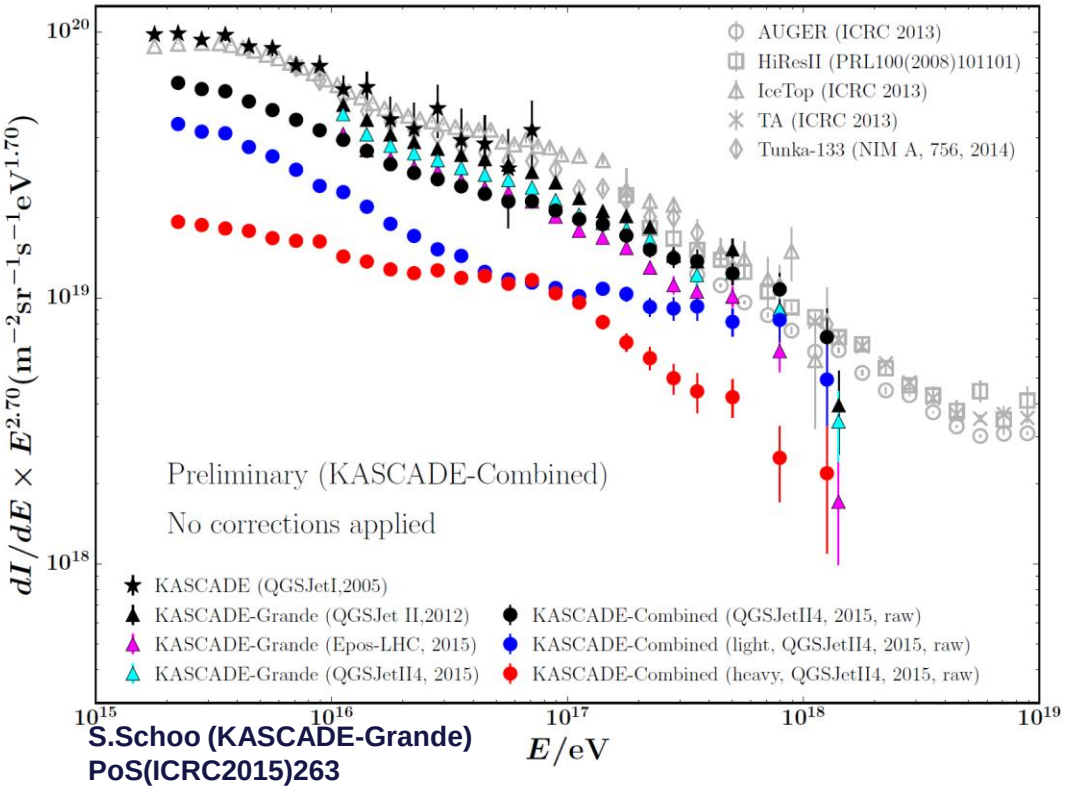


Post-LHC models are presently under investigation
EPOS-LHC looks a bit more promising

Juan Carlos Arteaga

Conclusion combined analysis:

All particle, light and heavy spectra for 3 orders of magnitude



S.Schoo (KASCADE-Grande)
PoS(ICRC2015)263

Paper in preparation
Analysis by Sven Schoo

- **Structures of spectra confirmed**
- **H4a model probably not far away from real composition**
- **Models still do not agree to each other and to data**
- **Light component seems to agree better than heavy**
- **Problem probably in the muons (known due to special selection)**
- **Around 10^{15} eV still (again) no clear picture**

Lessons learned from the >25-years KASCADE facility

It is essential to provide:

- spectra of individual mass groups!!
- multi-parameter EAS measurements to validate hadronic interaction models
- multi-messenger detection (need muons!!?)
- high statistics in a large energy range
  (mainly for composition dependent anisotropy studies)
- the right observation altitude
- room for R&D studies for future, improved technologies
- outreach and public data access

KASCADE-Grande Collaboration

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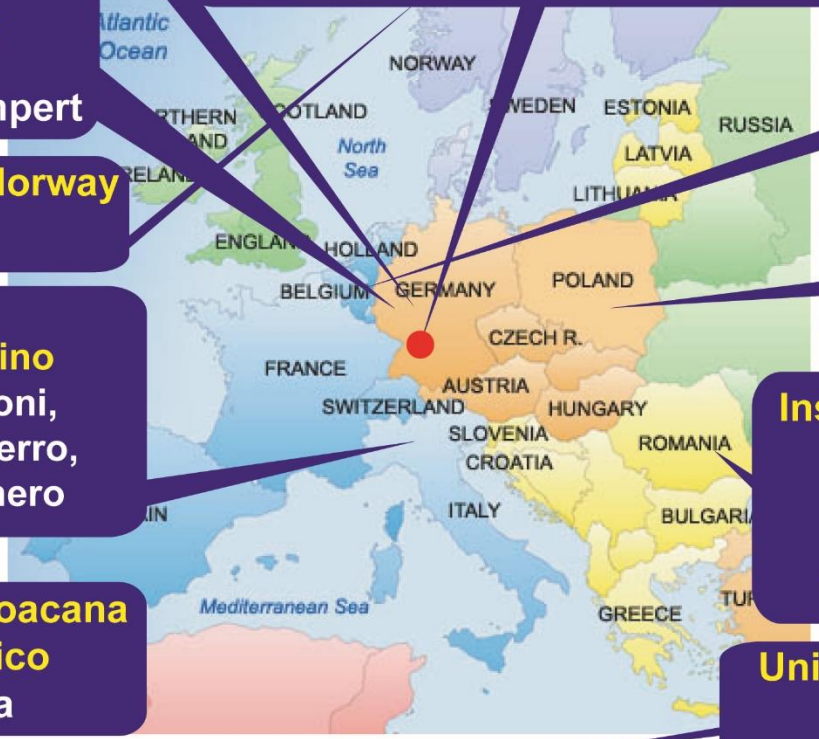

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KASCADE:
Contribution to most important question!

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Still better understanding of extensive air showers by improved hadronic interaction models are needed to answer this question

Hopefully not another 100 years (since V.Hess) or even 25 years (since KASCADE) needed to finally answer this question